
This is an electronic reprint of the original article.

This reprint may differ from the original in pagination and typographic detail.

Lister, M. L.; Aller, M. F.; Aller, H. D.; Homan, D. C.; Kellermann, K. I.; Kovalev, Y. Y.; Pushkarev, A. B.; Richards, J. L.; Ros, E.; Savolainen, T.

MOJAVE

Published in:
Astronomical Journal

DOI:
[10.3847/0004-6256/152/1/12](https://doi.org/10.3847/0004-6256/152/1/12)

Published: 01/07/2016

Document Version
Publisher's PDF, also known as Version of record

Please cite the original version:

Lister, M. L., Aller, M. F., Aller, H. D., Homan, D. C., Kellermann, K. I., Kovalev, Y. Y., Pushkarev, A. B., Richards, J. L., Ros, E., & Savolainen, T. (2016). MOJAVE: XIII. PARSEC-SCALE AGN JET KINEMATICS ANALYSIS BASED on 19 YEARS of VLBA OBSERVATIONS at 15 GHz. *Astronomical Journal*, 152(1), 1-16. Article 12. <https://doi.org/10.3847/0004-6256/152/1/12>



MOJAVE. XIII. PARSEC-SCALE AGN JET KINEMATICS ANALYSIS BASED ON 19 YEARS OF VLBA OBSERVATIONS AT 15 GHz

M. L. LISTER¹, M. F. ALLER², H. D. ALLER², D. C. HOMAN³, K. I. KELLERMANN⁴, Y. Y. KOVALEV^{5,6}, A. B. PUSHKAREV^{5,7}, J. L. RICHARDS¹, E. ROS^{6,8,9}, AND T. SAVOLAINEN^{6,10}

¹ Department of Physics and Astronomy, Purdue University, 525 Northwestern Avenue, West Lafayette, IN 47907, USA; mlister@purdue.edu

² Department of Astronomy, University of Michigan, 311 West Hall, 1085 S. University Avenue, Ann Arbor, MI:48109, USA

³ Department of Physics, Denison University, Granville, OH 43023, USA

⁴ National Radio Astronomy Observatory, 520 Edgemont Road, Charlottesville, VA 22903, USA

⁵ Astro Space Center of Lebedev Physical Institute, Profsoyuznaya 84/32, 117997 Moscow, Russia

⁶ Max-Planck-Institut für Radioastronomie, Auf dem Hügel 69, D-53121 Bonn, Germany

⁷ Crimean Astrophysical Observatory, 98409 Nauchny, Crimea, Russia

⁸ Observatori Astronòmic, Universitat de València, Parc Científic, C. Catedrático José Beltrán 2, E-46980 Paterna, València, Spain

⁹ Departament d'Astronomia i Astrofísica, Universitat de València, C. Dr. Moliner 50, E-46100 Burjassot, València, Spain

¹⁰ Aalto University Metsähovi Radio Observatory, Metsähovintie 114, FI-02540 Kylmäla, Finland

Received 2016 January 18; accepted 2016 March 9; published 2016 June 20

ABSTRACT

We present 1625 new 15 GHz (2 cm) VLBA images of 295 jets associated with active galactic nuclei (AGNs) from the MOJAVE and 2 cm VLBA surveys, spanning observations between 1994 August 31 and 2013 August 20. For 274 AGNs with at least 5 VLBA epochs, we have analyzed the kinematics of 961 individual bright features in their parsec-scale jets. A total of 122 of these jets have not been previously analyzed by the MOJAVE program. In the case of 451 jet features that had at least 10 epochs, we also examined their kinematics for possible accelerations. At least half of the well-sampled features have non-radial and/or accelerating trajectories, indicating that non-ballistic motion is common in AGN jets. Since it is impossible to extrapolate any accelerations that occurred before our monitoring period, we could only determine reliable ejection dates for $\sim 24\%$ of those features that had significant proper motions. The distribution of maximum apparent jet speeds in all 295 AGNs measured by our program to date is peaked below $5c$, with very few jets with apparent speeds above $30c$. The fastest speed in our survey is $\sim 50c$, measured in the jet of the quasar PKS 0805–07, and is indicative of a maximum jet Lorentz factor of ~ 50 in the parent population. An envelope in the maximum jet speed versus redshift distribution of our sample provides additional evidence of this upper limit to the speeds of radio-emitting regions in parsec-scale AGN jets. The *Fermi*-LAT-detected gamma-ray AGNs in our sample have, on average, higher jet speeds than non-LAT-detected AGNs, indicating a strong correlation between parsec-scale jet speed and the gamma-ray Doppler boosting factor. We have identified 11 moderate-redshift ($z < 0.35$) AGNs with fast apparent speeds ($> 10c$) that are strong candidates for future TeV gamma-ray detection. Of the five gamma-ray loud narrow-lined Seyfert I AGNs in our sample, three show highly superluminal jet motions, while the others have sub-luminal speeds. This indicates that some narrow-lined Seyfert I AGNs possess powerful jets with Lorentz factors in excess of 10, and viewing angles less than 10° , consistent with those of typical BL Lac objects and flat-spectrum radio quasars.

Key words: BL Lacertae objects: general – galaxies: active – galaxies: jets – quasars: general – radio continuum: galaxies

Supporting material: animations, figure set, machine-readable tables

1. INTRODUCTION

Very long baseline interferometry (VLBI) provides the highest resolution images in astronomy and is invaluable for studying the structure and evolution of non-thermal radio sources. The powerful jetted outflows associated with active galactic nuclei (AGNs) are a natural target for VLBI, since they are exceedingly compact, and dominate the extragalactic sky at radio and high-energy wavelengths. Because of their relativistic flow speeds, AGN jets are affected by beaming and aberration effects that can significantly enhance their apparent luminosity and variability, and create the illusion of superluminal motion in the sky plane (Blandford et al. 1977). AGNs of the rare blazar class represent cases where a high Lorentz factor jet is pointed nearly directly at us, thereby maximizing these relativistic Doppler effects.

Starting with the 2 cm VLBA survey in 1994 (Kellermann et al. 1998), and continuing with the MOJAVE (Monitoring of Jets in Active Galactic Nuclei with VLBA Experiments) program (Lister et al. 2009a), we have carried out multi-epoch

VLBA observations of several hundred of the brightest, most compact radio sources in the northern sky. These are predominantly blazars, due to the selection biases associated with relativistic beaming (Blandford and Königl 1979).

In this paper, we present 1625 VLBA 15 GHz contour maps of 295 AGNs for epochs between 1994 August 31 and 2013 August 20 that have not previously appeared in any paper from the MOJAVE (Lister & Homan 2005; Lister et al. 2009a, 2013) or 2 cm VLBA surveys (Kellermann et al. 1998; Zensus et al. 2002). These AGNs (see Table 1) are from one or more of the following: the MOJAVE low-luminosity sample (Lister et al. 2013), the complete flux density-limited MOJAVE 1.5 Jy sample (Lister et al. 2015), the VLBA 2 cm survey (Kellermann et al. 1998), the third EGRET gamma-ray catalog (Hartman et al. 1999), or the 3FGL *Fermi* gamma-ray catalog (Acero et al. 2015). Also included are some AGNs that were originally candidates for these samples, but did not meet the final selection criteria.

Table 1
AGN Properties

B1950 (1)	J2000 (2)	Alias (3)	Gamma-ray Assoc. (4)	Opt. (5)	1.5 Jy (6)	z (7)	Reference (8)
0003+380 ^a	J0005+3820	S4 0003+38	3FGL J0006.4+3825	Q	...	0.229	Schramm et al. (1994)
0003-066 ^a	J0006-0623	NRAO 005	...	B	Y	0.3467	Jones et al. (2005)
0006+061 ^a	J0009+0628	CRATES J0009+0628	3FGL J0009.1+0630 ^c	B	Rau et al. (2012)
0007+106 ^a	J0010+1058	III Zw 2	...	G	Y	0.0893	Sargent (1970)
0010+405 ^a	J0013+4051	4C +40.01	...	Q	...	0.256	Thompson et al. (1992)
0015-054 ^a	J0017-0512	PMN J0017-0512	3FGL J0017.6-0512	Q	...	0.226	Shaw et al. (2012)
0016+731 ^a	J0019+7327	S5 0016+73	...	Q	Y	1.781	Lawrence et al. (1986)
0027+056 ^a	J0029+0554	PKS 0027+056	...	Q	...	1.317	Schneider et al. (1999)
0026+346 ^a	J0029+3456	B2 0026+34	...	G	...	0.517	Zensus et al. (2002)
0035+413 ^a	J0038+4137	B3 0035+413	...	Q	...	1.353	Stickel & Kuhr (1993)

Notes. Columns are as follows: (1) B1950 name, (2) J2000 name, (3) other name, (4) gamma-ray association name, (5) optical classification, where B = BL Lac, Q = quasar, G = radio galaxy, N = narrow-lined Seyfert 1, and U = unidentified, (6) MOJAVE 1.5 Jy sample membership flag, (7) redshift, (8) reference for redshift and/or optical classification.

^a Jet kinematics analyzed in this paper, based on data up to 2013 August 20.

^b Jet kinematics analyzed by Lister et al. (2013), based on data up to 2011 May 1.

^c Known TeV emitter (<http://tevcat.uchicago.edu>).

^d *Fermi*-LAT detection reported by Carpenter et al. (2014).

^e One of two AGN associations listed for this 3FGL source in Ackermann et al. (2015; see Appendix A).

(This table is available in machine-readable form.)

A major aspect of our program is to analyze these multi-epoch VLBA observations in order to investigate the evolution of parsec-scale AGN jets, which we have reported on in a series of papers (Kellermann et al. 2004; Homan et al. 2009; Lister et al. 2009b, 2013; Homan et al. 2015). Most AGN jets have bright, compact radio-emitting features that move outward at apparent superluminal speeds. These features often change both their speed and direction, suggesting that their kinematics are strongly affected by both hydrodynamic and MHD plasma effects (e.g., Homan et al. 2003; Cohen et al. 2015). In our most recent analysis (Lister et al. 2013; Homan et al. 2015), we studied the kinematics of 200 AGN jets, for which at least 5 VLBA epochs were obtained between 1994 August 31 and 2011 May 1. These AGNs were drawn from the MOJAVE 1.5 Jy radio flux density-limited and 1FM *Fermi* gamma-ray selected samples, described by Lister et al. (2015) and Lister et al. (2011), respectively. Here we present time-lapse movies and new kinematics analyses of 173 AGNs drawn from these two samples, and 101 AGNs drawn from the other surveys described in the previous paragraph, using VLBA data obtained between 1994 August 31 and 2013 August 20. A total of 122 of these jets have not previously had their kinematics analyzed by the MOJAVE program. We have excluded from our kinematics analysis 48 MOJAVE AGNs we previously analyzed in Lister et al. (2013), since we are currently gathering additional VLBA epochs on them.

The overall layout of the paper is as follows. In Section 2, we describe our observational data and our method of modeling the individual jet features, and, in Section 3, we discuss their kinematic properties and overall trends in the data. We summarize our findings in Section 4. We adopt a cosmology with $\Omega_m = 0.27$, $\Omega_\Lambda = 0.73$, and $H_0 = 71 \text{ km s}^{-1} \text{ Mpc}^{-1}$ (Komatsu et al. 2009). We refer to the AGNs throughout using either B1950 nomenclature or commonly used aliases, as listed in Table 1.

2. OBSERVATIONAL DATA AND ANALYSIS

2.1. Contour Maps

We processed the VLBA interferometric visibility data and produced contour maps (Figure 1) using the AIPS and Difmap (Shepherd 1997) software packages, using the methods described by Lister et al. (2009a, 2013). The raw correlated data from the MOJAVE program are available immediately after correlation from the NRAO data archive, and our processed data and maps are publicly available on our project website¹¹ within a few weeks after correlation. We list the parameters of the contour maps in Table 2. The FWHM dimensions of the VLBA interferometric restoring beam at 15 GHz vary with declination, but are on the order of 1 mas (N-S) \times 0.5 mas (E-W), which corresponds to a linear resolution of a few parsecs (projected) at the median redshift ($z \simeq 1$) of our sample. Column 4 of Table 2 lists the VLBA project code for each observation, along with an indicator of whether it is from the MOJAVE program, the VLBA 2 cm Survey, or the NRAO data archive. For the latter, we considered only archival epochs with at least four individual scans that spanned a reasonably wide range of hour angle, and that included at least eight VLBA antennas. The VLBA 2 cm Survey observations (1994–2002) consisted of approximately one-hour integrations on each AGN, broken up into several scans separated in hour angle to improve the interferometric coverage. A similar observing method was used in the full polarization MOJAVE observations from 2002 May to 2007 September (VLBA codes BL111, BL123, BL137, and BL149) and is described by Lister & Homan (2005). During 2006 (VLBA code BL137), the 15 GHz integration times were shortened by a factor of approximately three to accommodate interleaved scans at three other observing frequencies (8.1, 8.4, 12.1 GHz). The latter were published by Hovatta et al. (2012), Pushkarev et al. (2012), and Hovatta et al. (2014). The

¹¹ <http://www.astro.purdue.edu/MOJAVE>

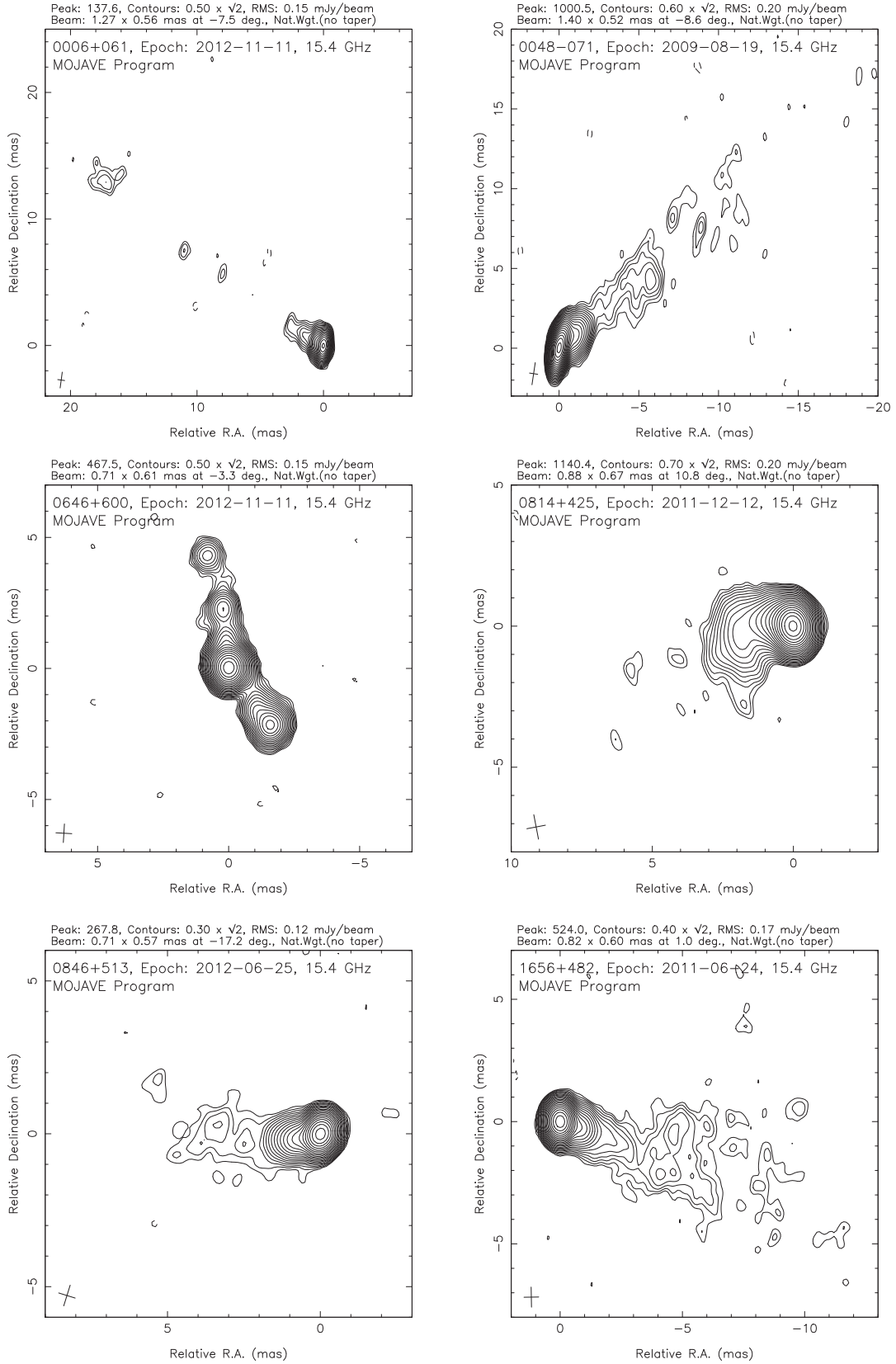


Figure 1. Naturally weighted 15 GHz total intensity VLBA contour images of individual epoch observations of the MOJAVE AGN sample. The contours are in successive powers of $\sqrt{2}$ times the base contour level in mJy per beam, as listed in Table 2 and at the top of each panel. The FWHM restoring beam dimensions are indicated as a cross in the lower left corner. Because of self-calibration, in some cases the origin may be coincident with the brightest feature in the image, rather than the putative core feature listed in Table 3.

(The complete figure set (64 images) is available.)

Table 2
Summary of 15 GHz Image Parameters

Source (1)	Alias (2)	Epoch (3)	VLBA Code (4)	Freq. (GHz) (5)	B_{maj} (mas) (6)	B_{min} (mas) (7)	B_{pa} ($^{\circ}$) (8)	I_{tot} (Jy) (9)	rms (mJy bm^{-1}) (10)	I_{base} (mJy bm^{-1}) (11)	Fig. Num. (12)
0003+380	S4 0003+38	2011 Jun 6	BL149DJ ^a	15.4	0.91	0.54	−10	0.605	0.18	0.50	1.1
		2013 Aug 12	BL178BH ^a	15.4	0.84	0.53	−4	0.667	0.20	0.50	1.2
0003−066	NRAO 005	2011 Jun 24	BL149DL ^a	15.4	1.42	0.54	−5	2.076	0.18	1.30	1.3
		2012 Mar 4	BL178AH ^a	15.4	1.31	0.53	−3	2.132	0.17	0.60	1.4
		2012 Nov 2	BL178AR ^a	15.4	1.37	0.53	−7	2.273	0.16	0.60	1.5
0006+061	CRATES J0009+0628	2011 Dec 29	BL178AD ^a	15.4	1.32	0.65	8	0.209	0.22	0.60	1.6
		2012 Jul 12	BL178AM ^a	15.4	1.28	0.52	−13	0.200	0.25	0.70	1.7
		2012 Nov 11	BL178AS ^a	15.4	1.27	0.56	−8	0.182	0.15	0.50	1.8
		2012 Dec 23	BL178AX ^a	15.4	1.34	0.53	−3	0.184	0.20	0.60	1.9
		2013 Jun 2	BL178BD ^a	15.4	1.32	0.56	−1	0.162	0.23	0.70	1.10
0007+106	III Zw 2	1998 Feb 16	BF039 ^c	15.4	1.14	0.51	−8	0.784	0.16	0.50	1.11
		1998 Jun 13	BF039B ^c	15.4	1.13	0.53	−5	0.982	0.15	0.40	1.12
		1998 Sep 14	BG044 ^c	15.4	1.11	0.49	−4	1.202	0.18	0.70	1.13
		2011 May 26	BL149DI ^a	15.4	1.18	0.54	−5	0.169	0.19	0.60	1.14

Notes. Columns are as follows: (1) B1950 name, (2) other name, (3) date of VLBA observation, (4) VLBA experiment code, (5) observing frequency (GHz), (6) FWHM major axis of restoring beam (milliarcseconds), (7) FWHM minor axis of restoring beam (milliarcseconds), (8) position angle of major axis of restoring beam (degrees), (9) total I flux density (Jy), (10) rms noise level of image (mJy per beam), (11) lowest I contour (mJy per beam), (12) figure number.

^a Full polarization MOJAVE epoch.

^b 2 cm VLBA Survey epoch.

^c NRAO archive epoch.

(This table is available in its entirety in machine-readable form.)

MOJAVE and 2 cm Survey observations were originally recorded at a data rate of 128 Mbps, which was increased to 256 Mbps for the epochs from 2007 July 3 to 2008 September 12 inclusive, and 512 Mbps thereafter. Beginning with the 2007 January 6 epoch, we increased the number of AGNs observed in each 24 hr MOJAVE session from 18 to 25 AGNs, and to 30 AGNs on 2009 February 25.

2.2. Time-lapse Movies

In Figures 2 and 3, we show linearly interpolated time-lapse MPEG movies of the multi-epoch VLBA images for two selected AGNs (4C +67.05 and PKS 2345−16). Movies for the other AGNs can be found on the MOJAVE website. We constructed the movies using a two-point linear interpolation across each successive epoch, treating each map pixel independently. Prior to interpolation, we restored all of the epoch maps to a scale of 0.05 mas per pixel using identical median beam dimensions that were based on the full set of naturally weighted VLBA epochs available for that AGN. We did not interpolate across any time gaps larger than four years, leaving these periods instead as blank frames in the movie. The false-color corresponds to radio flux density in units of jansky per beam (indicated by scale bar on the right side of the frames). One year of calendar time corresponds to 2.5 s of run time in the movies.

2.3. Gaussian Modeling

We modeled the (u, v) visibility data in Difmap using a set of Gaussian features, which we list in Table 3. In some instances, it was not possible to robustly cross-identify the same features in a jet from one epoch to the next. We indicate those features with robust cross-identifications across at least five epochs in column 10 of Table 3. For the non-robust features, we caution that the assignment of the same identification number across

epochs does not necessarily indicate a reliable cross-identification.

Based on previous analysis (Lister et al. 2009b), we estimate the typical uncertainties in the Gaussian centroid positions to be $\sim 20\%$ of the FWHM beam dimensions. For isolated bright and compact features, the positional errors are smaller by approximately a factor of two. We estimate the formal errors on the feature sizes to be roughly twice the positional error, according to Fomalont (1999). The flux density accuracies are approximately 5% (see Appendix A of Homan et al. 2002), but can be significantly larger for features located very close to one another. Also, at some epochs which had antenna dropouts, the fit errors of some features are much larger. We do not use the latter in our kinematics analysis and indicate them by flags in Table 3.

2.4. Jet Kinematics Analysis

As in Lister et al. (2009b, 2013) and Homan et al. (2015, 2009), we analyzed the kinematics of the robust Gaussian jet features in our sample using two methods. In the first method, we assumed a non-accelerating, two-dimensional (R.A. and decl.) vector fit to the feature position over time, referenced to the core feature (which we presumed to be stationary). The latter represents the region in our VLBA images near the base of the jet where the emission becomes optically thick at 15 GHz. For the features that had measurements at 10 or more epochs, our second method employed a constant acceleration model that yielded kinematic fit quantities at a (midpoint) reference date located exactly halfway between the first and last VLBA observation dates for that particular AGN (see column 9 of Table 5). The results of these analyses are listed separately in Tables 4 and 5, respectively. For completeness, these tables also include kinematic fit parameters for 48 AGNs (indicated by flags in column 1) that we analyzed in Lister et al. (2013), but not in this paper. We note that for features that show

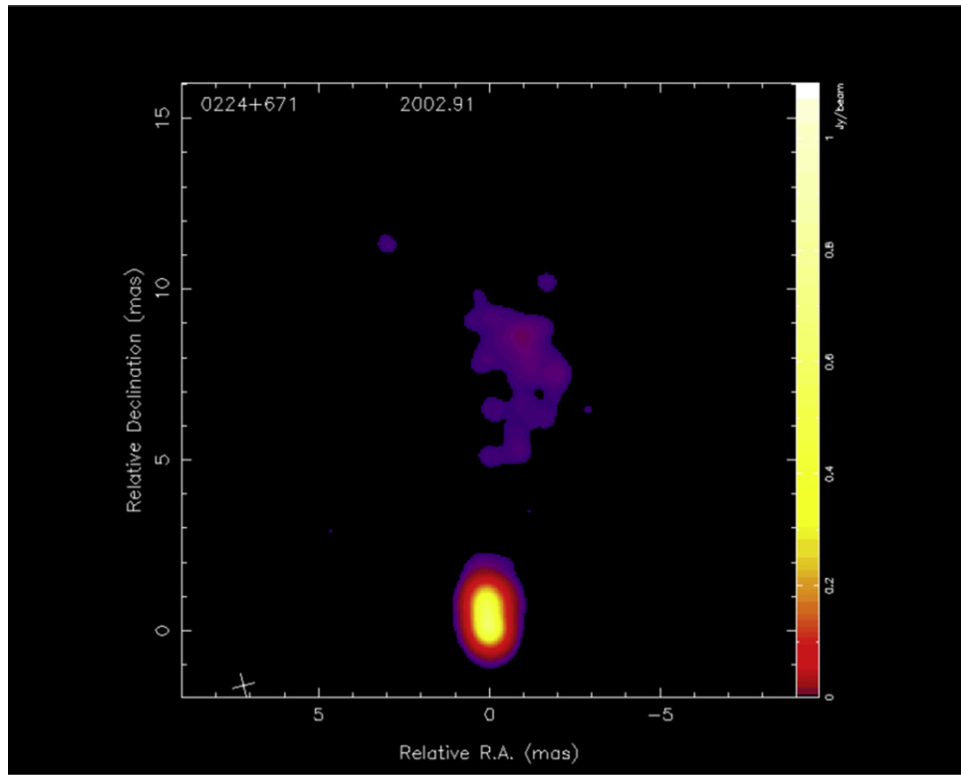


Figure 2. Linearly interpolated time-lapse movie made from the multi-epoch VLBA images for 4C +67.05. Each epoch is restored with a median beam, whose FWHM dimensions are indicated by the cross in the lower left corner. The false-color corresponds to radio flux density in units of Jy per beam (indicated by scale bar on the right side of the frames). One year of calendar time corresponds to 2.5 s in the movies.

(An animation of this figure is available.)

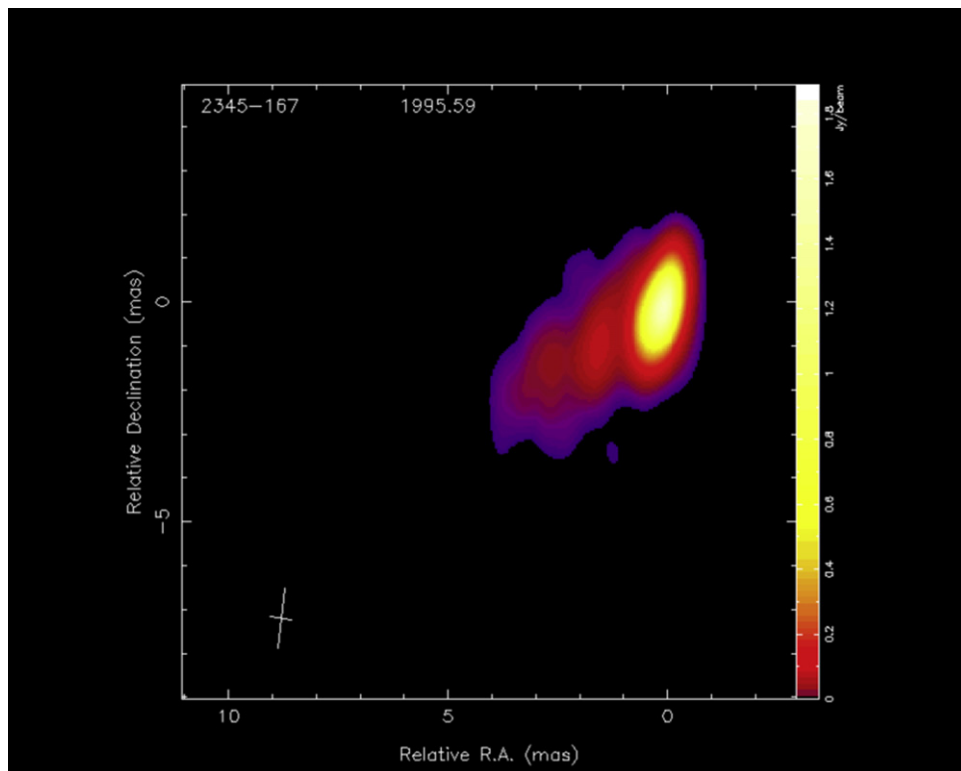


Figure 3. Same as Figure 2, for the quasar PKS 2345-16.

(An animation of this figure is available.)

Table 3
Fitted Jet Features

Source (1)	I.D. (2)	Epoch (3)	I (Jy) (4)	r (mas) (5)	P.A. ($^{\circ}$) (6)	Maj. (mas) (7)	Ratio (8)	Maj. P.A. ($^{\circ}$) (9)	Robust? (10)
0003+380	0	2006 Mar 9	0.489	0.04	290.7	0.23	0.33	292	Y
	1	...	0.007	3.98	121.8	0.72	1.00	...	Y
	2	...	0.042	1.25	110.5	0.51	1.00	...	Y
	6	...	0.104	0.28	114.6	0.27	1.00	...	Y
	7	...	0.003	2.31	119.3	N
	0	2006 Dec 1	0.320	0.10	308.1	0.25	0.29	295	Y
	1	...	0.005	3.65	120.8	1.63	1.00	...	Y
	2	...	0.021	1.56	111.0	0.25	1.00	...	Y
	5	...	0.023	0.75	116.2	0.32	1.00	...	Y
	6	...	0.145	0.45	116.3	0.05	1.00	...	Y

Notes. Columns are as follows: (1) B1950 name, (2) feature identification number (zero indicates core feature), (3) observation epoch, (4) flux density in jansky, (5) position offset from the core feature (or map center for the core feature entries) in milliarcseconds, (6) position angle with respect to the core feature (or map center for the core feature entries) in degrees, (7) FWHM major axis of fitted Gaussian in milliarcseconds, (8) axial ratio of fitted Gaussian, (9) major axis position angle of fitted Gaussian in degrees, (10) robust feature flag.

^a Individual feature epoch not used in kinematic fits.

(This table is available in its entirety in machine-readable form.)

Table 4
Vector Motion Fit Properties of Jet Features

Source (1)	I.D. (2)	N (3)	$\langle S \rangle$ (mJy) (4)	$\langle R \rangle$ (mas) (5)	$\langle d_{\text{proj}} \rangle$ (pc) (6)	$\langle \vartheta \rangle$ (deg) (7)	ϕ (deg) (8)	$ \langle \vartheta \rangle - \phi $ (deg) (9)	μ (μasy^{-1}) (10)	β_{app} (c) (11)	t_{ej} (12)	t_{α} (13)	t_{δ} (14)
0003+380	1	9	5	4.24	15.40	120.5	97 ± 16	24 ± 16	157 ± 42	2.29 ± 0.61	...	1985.09	1888.98
	2	7	17	1.81	6.57	112.6	119.7 ± 2.6	7.1 ± 2.7	319 ± 22	4.64 ± 0.32	...	2001.68	2003.32
	4	6	17	1.25	4.54	114.8	208 ± 12	93 ± 12^a	36.3 ± 9.2	0.53 ± 0.13	...	2075.66	1992.93
	5	9	39	0.76	2.77	117.5	331 ± 119	146 ± 119	4.8 ± 7.2^b	0.07 ± 0.10	...	2300.95	2092.52
	6	10	98	0.39	1.43	115.4	335 ± 41	141 ± 41	12.7 ± 8.3^b	0.19 ± 0.12	...	2073.84	2023.64
	2	5	222	1.05	5.12	322.9	226.3 ± 4.8	96.6 ± 5.0^a	191 ± 15	4.09 ± 0.33	...	1993.56	2004.48
0003-066	3	9	119	2.82	13.73	296.9	284.8 ± 4.5	12.1 ± 4.6	250 ± 39	5.36 ± 0.83	...	1989.50	1979.99
	4 ^c	26	119	6.61	32.23	285.6	283.6 ± 8.1	1.9 ± 8.1	41 ± 13	0.87 ± 0.28	...	1845.03	1820.94
	5	6	1031	0.73	3.58	14.5	343.1 ± 3.1	31.4 ± 3.1^a	100 ± 16	2.15 ± 0.35	...	2011.53	1997.84
	6 ^c	10	97	1.01	4.92	290.2	210 ± 13	81 ± 13^a	55 ± 17	1.18 ± 0.37	...	1969.25	2011.24

Notes. Columns are as follows: (1) B1950 name, (2) feature number, (3) number of fitted epochs, (4) mean flux density at 15 GHz in millijansky, (5) mean distance from core feature in mas, (6) mean projected distance from core feature in parsec, (7) mean position angle with respect to the core feature in degrees, (8) position angle of velocity vector in degrees, (9) offset between mean position angle and velocity vector position angle in degrees, (10) angular proper motion in mas per year, (11) fitted speed in units of the speed of light, (12) estimated epoch of origin, (13) fitted epoch of origin in right ascension direction, (14) fitted epoch of origin in declination direction.

A question mark indicates a feature whose motion is not consistent with outward, radial motion, but for which the possibility of inward motion and its degree of non-radialness are uncertain.

^a Feature has significant non-radial motion.

^b Feature has slow pattern speed.

^c Feature has significant accelerated motion (see Table 5 for acceleration fit parameters).

^d Feature has significant inward motion.

^e Fit parameters for this source are from Lister et al. (2013).

(This table is available in its entirety in machine-readable form.)

significant ($\geq 3\sigma$) accelerations, the acceleration-fit speed and other parameters in Table 5 provide a better description of their motions, hence for these features we will use the acceleration fit quantities in subsequent statistical analyses and discussion of their kinematics.

In the left-hand panels of Figure 4, we show the positions and kinematic motion fits to the individual features on the sky, as well as a 15 GHz VLBA contour image of the jet at the observation epoch that is closest to the (midpoint) reference

date. The orange box delimits the zoomed region displayed in the right-hand panels, with the cross-hairs indicating the feature's position at this epoch.

We note that many of the jet features in our sample display non-radial motion, i.e., the fitted velocity vector at the reference epoch does not point back to the core feature. This is indicative of accelerated, non-ballistic motion. In some cases, an extrapolation of our simple constant acceleration fit shows the feature passing through the core feature at the ejection

Table 5
Acceleration Fit Properties of Jet Features

Source (1)	I.D. (2)	ϕ (deg) (3)	$ \langle\vartheta\rangle - \phi $ (deg) (4)	μ ($\mu\text{as yr}^{-1}$) (5)	β_{app} (c) (6)	$\dot{\mu}_{\perp}$ ($\mu\text{as yr}^{-2}$) (7)	$\dot{\mu}_{\parallel}$ ($\mu\text{as yr}^{-2}$) (8)	t_{mid} (9)	t_{α} (10)	t_{δ} (11)
0003+380	6	333 ± 40	142 ± 40	13.4 ± 8.5	0.20 ± 0.12	$-4 \pm 11 (\pm 8.3)$	$9.0 \pm 9.2 (\pm 7.6)$	2009.90	2071.12	2024.51
0003–066	4 ^a	270.0 ± 4.1	15.6 ± 4.1^b	50.4 ± 5.3	1.08 ± 0.11	$8.5 \pm 2.5 (\pm 1.6)$	$-27.4 \pm 2.4 (\pm 2.3)$	2004.83	1874.35	503767
	6 ^a	211.3 ± 8.8	78.9 ± 8.9^b	54 ± 11	1.16 ± 0.24	$54 \pm 13 (\pm 11)$	$-37 \pm 17 (\pm 15)$	2003.78	1971.75	2009.60
	8 ^a	290.7 ± 1.6	3.5 ± 1.7	330.4 ± 9.8	7.08 ± 0.21	$-19 \pm 12 (\pm 12)$	$-64 \pm 12 (\pm 12)$	2009.93	2002.02	2000.32
	9	295.2 ± 4.1	7.5 ± 4.2	278 ± 20	5.96 ± 0.42	$9 \pm 36 (\pm 36)$	$-99 \pm 35 (\pm 35)$	2009.24	2002.22	2004.34
0010+405	1	340.7 ± 4.3	11.9 ± 4.3	432 ± 42	6.99 ± 0.68	$11 \pm 54 (\pm 53)$	$-43 \pm 69 (\pm 69)$	2008.52	1978.76	1991.39
	2	359 ± 237	31 ± 237	3 ± 17	0.04 ± 0.28	$2 \pm 16 (\pm 16)$	$-3 \pm 26 (\pm 26)$	2008.87	-37937	1427.76
	3	138 ± 94	170 ± 94	2.6 ± 4.6	0.043 ± 0.074	$-4.3 \pm 9.6 (\pm 6.6)$	$5.4 \pm 9.8 (\pm 7.0)$	2008.87	2290.10	2404.54
	4	114 ± 190	146 ± 190	1.4 ± 3.5	0.023 ± 0.057	$-2.1 \pm 8.8 (\pm 7.2)$	$-4.7 \pm 7.5 (\pm 5.4)$	2008.87	2195.70	2676.67
0016+731	1 ^a	163.2 ± 2.2	22.8 ± 2.4^b	106.2 ± 4.4	8.23 ± 0.34	$10.2 \pm 1.8 (\pm 1.8)$	$8.8 \pm 2.0 (\pm 1.9)$	2002.70	1975.08	1994.50

Notes. Columns are as follows: (1) B1950 name, (2) feature number, (3) proper motion position angle in degrees, (4) offset between mean position angle and velocity vector position angle in degrees, (5) angular proper motion in mas per year, (6) fitted speed in units of the speed of light, (7) angular acceleration perpendicular to velocity direction in mas per year per year, with the error excluding the direction uncertainty in parentheses, (8) angular acceleration parallel to velocity direction in mas per year per year, with the error excluding the direction uncertainty in parentheses, (9) date of reference (middle) epoch used for fit, (10) fitted epoch of origin in right ascension direction, (11) fitted epoch of origin in declination direction.

A question mark indicates a feature whose motion is not consistent with outward, radial motion, but for which the possibility of inward motion and its degree of non-radialness are uncertain.

^a Feature shows significant accelerated motion.

^b Feature shows significant non-radial motion.

^c Feature shows significant inward motion.

^d Fit parameters for this source are from Lister et al. (2013).

(This table is available in its entirety in machine-readable form.)

epoch within the errors, as expected. However, in many other cases, the acceleration(s) are either complex, occurred before our monitoring period, or took place close to the core below our angular resolution level. These unknown accelerations strictly limit the number of features for which we can derive reliable ejection epochs, and can cause an apparent inflection in the fit line at minimum separation (e.g., feature id = 5 of 0736 + 017) in the plots of radial separation from the core versus time (see Figure 5).

We examined each feature with significant proper motion ($\mu \geq 3\sigma_{\mu}$) for non-radial motion by comparing the mean position angle of the feature $\langle\vartheta\rangle$ with its proper motion vector direction ϕ . We flagged any feature for which the angular offset $|\langle\vartheta\rangle - \phi|$ was $\geq 3\sigma$ from both 0° and 180° as non-radial, and inward if the offset was significantly greater than 90° . Of the 714 newly analyzed features with statistically significant ($\geq 3\sigma$) proper motion, 228 (32%) exhibit significant non-radial motion, while only 25 (4%) are flagged as inward (see Table 4). The latter can be the result of curved trajectories that are approaching the line of sight, internal brightness changes in a large, diffuse jet feature, or backward pattern speeds not associated with the flow. We discuss these rare individual inward motion cases in the Appendix.

We calculated ejection times (epoch of zero core separation) for the non-accelerating, non-inward-moving features with significant motion by taking an error-based weighted average of the extrapolated emergence epochs in the right ascension (t_{α}) and declination (t_{δ}) directions (columns 13 and 14 of Table 4). The individual values of t_{α} and t_{δ} describe the fits, but do not necessarily have physical meaning, depending on the proper motion vector direction, the statistical significance of the fitted motion, and/or accelerations that may be present. We include them in the motion tables for completeness only. We used the

method described in Lister et al. (2013) to calculate the errors on the ejection times.

Due to the nature of our proper motion fits, which naturally include the possibility of non-ballistic motion, we did not estimate ejection times for any features where we could not confidently extrapolate their motion to the core. Jet features for which we estimated ejection epochs had the following properties: (1) significant motion ($\mu \geq 3\sigma_{\mu}$), (2) no significant acceleration, (3) a velocity vector within 15° of the outward radial direction to high confidence, i.e., $|\langle\vartheta\rangle - \phi| + 2\sigma \leq 15^\circ$, (4) an extrapolated position at the ejection epoch no more than 0.2 mas from the core, and (5) a fitted ejection epoch that differed by no more than 0.5 years from that given by a simple one-dimensional (radial) motion fit.

We list the parameters of our acceleration fits in Table 5, where we have resolved the acceleration terms $\dot{\mu}_{\perp}$ and $\dot{\mu}_{\parallel}$ in directions perpendicular and parallel, respectively, to the mean angular velocity direction ϕ (columns 7 and 8). We quote two uncertainty values in these acceleration terms. The value in parentheses is the statistical uncertainty due to the acceleration fit itself and is the appropriate quantity for determining if a significant acceleration occurred relative to the velocity vector direction. The total uncertainty value incorporates an estimate of the effect of a possibly poorly known velocity direction when assigning these accelerations the labels “parallel” or “perpendicular”. For example, the uncertainty on the parallel acceleration adds in quadrature the term $\sqrt{(\dot{\mu}_{\perp} d\phi)^2 + ((\dot{\mu}_{\parallel} d\phi) d\phi)^2}$ where $d\phi$ is the uncertainty in the velocity direction in radians, and takes on a maximum value of 1.0 in this expression. An analogous term is added in quadrature to the statistical uncertainty for the perpendicular acceleration. In our previous papers, we listed the acceleration fit uncertainty alone, and did not include any additional error

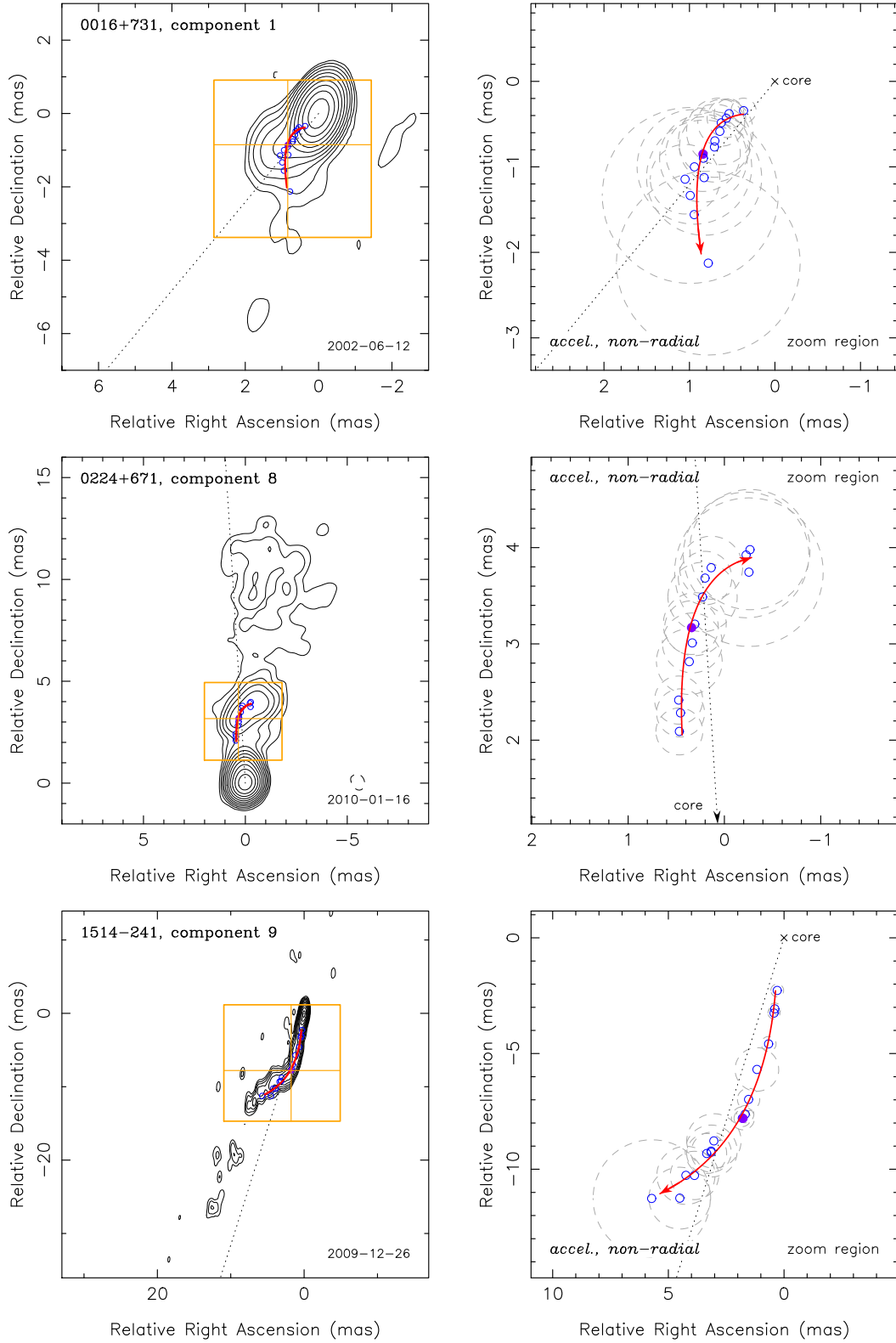


Figure 4. Vector motion fits and sky position plots of individual robust jet features in MOJAVE AGNs. Positions are relative to the core position. The left-hand panels show a 15 GHz VLBA contour image of the jet at the epoch listed in Table 5, which is closest to the (midpoint) reference epoch. The orange box delimits the zoomed region that is displayed in the right-hand panels. The feature's position at the image epoch is indicated by the orange cross-hairs. The dotted line connects the feature with the core feature and is plotted with the mean position angle $\langle \theta \rangle$. The position at the image epoch is shown by a filled violet circle while other epochs are plotted with unfilled blue circles. The red solid line indicates the vector fit (or accelerating fit, if there is significant acceleration) to the feature positions. The gray dashed circles/ellipses represent the FWHM sizes of the individual fitted Gaussian features. A red vector motion direction arrow is indicated in all of the plots, including features for which the angular speed is extremely slow and/or statistically insignificant.

(The complete figure set (64 images) is available.)

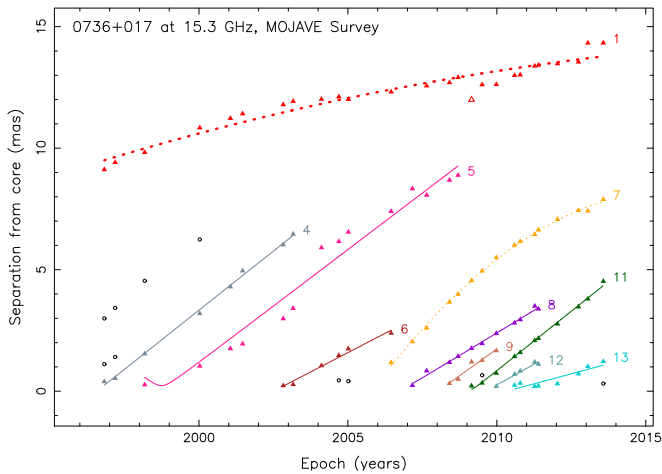


Figure 5. Plot of angular separation from the core vs. epoch for Gaussian jet features. The B1950 source name is given at the top left of each panel. Colored symbols indicate robust features for which kinematic fits were obtained (dotted and solid lines). The solid lines indicate vector motion fits to the data points assuming no acceleration, while the dotted lines indicate accelerated motion fits. Thick lines are used for features whose fitted motion is along a radial direction from the core, while thin lines indicate non-radial motions. Unfilled colored circles indicate individual data points that were not used in the kinematic fits, and unfilled black circles indicate non-robust features. The feature identification number is indicated next to the last epoch of each robust feature. (This is a figure stub; an extended version is available online.)

(The complete figure set (64 images) is available.)

due to uncertainty in the vector motion direction. Instead, as described in Homan et al. (2009, 2015), we limited our detailed acceleration analysis in those papers to jet features where the velocity direction was very well known and therefore the parallel and perpendicular assignments were unambiguous.

Of the 451 features in our current analysis with more than 10 epochs, 144 (32%) show significant accelerations, and an additional 87 show non-radial motions. Thus over half of the well-sampled jet features in our survey show evidence of non-ballistic motion. In agreement with our previous findings (Homan et al. 2009, 2015), apparent speed changes are larger and more prevalent than changes in direction, indicating that changes in Lorentz factor are more responsible for the observed accelerations than jet bends.

3. DISCUSSION

3.1. Jet Speed Distributions

We have calculated maximum and median speed statistics for all 274 AGN jets measured in this paper, and 48 AGNs in Lister et al. (2013), using the method described in the latter paper. For those AGNs that were analyzed in both papers, we used the kinematic results from the current paper. The combined set of 322 AGNs spans the full sky above declination -30° and all sources have at least ~ 0.1 Jy of VLBA correlated flux density at 15 GHz. Although the set is complete above 1.5 Jy (Lister et al. 2015) and at high γ -ray (time-averaged) fluxes (Lister et al. 2011), the selection is incomplete and non-uniform at lower radio flux densities, containing a sampling of gamma-ray-associated and unassociated AGNs, as well as lower-luminosity AGNs targeted specifically to probe the relation between jet speed and radio luminosity (Lister et al. 2013). Because of the >0.1 Jy VLBA limit at 15 GHz, the

combined set of AGNs has a substantial bias toward core-dominated, flat-spectrum radio sources, which are expected to be highly beamed. This is reflected in the optical classifications, which are dominated by blazars (i.e., flat-spectrum radio quasars: 73% and BL Lac objects: 18%). There are also 22 radio galaxies and 5 narrow-lined Seyfert I (NLSY1) galaxies, and 4 AGNs with no optically identified counterparts. A total of 243 AGNs have published gamma-ray associations, which are listed in column 4 of Table 1.

In Figure 6, we plot the speed distributions for the combined set of AGNs (left-hand panels), and for the radio flux density-limited MOJAVE 1.5 Jy sample only (right-hand panels). As discussed by Vermeulen & Cohen (1994) and Lister & Marscher (1997), the shape of the apparent speed distribution in a flux density-limited jet sample reflects the parent Lorentz factor distribution, which, in this case, indicates a moderately steep parent power law, with fast jets being relatively rare. A more detailed analysis of the parent population properties will be presented in a future paper. Because of Doppler bias, the fastest apparent speeds in a large flux density-limited sample should also be indicative of the maximum bulk flow Lorentz factors in the parent population. The fastest overall jet speed we have measured in our survey to date is feature id = 2 in the quasar PKS 0805–07, with $\mu = 503 \pm 62 \mu\text{as yr}^{-1}$ corresponding to $40c \pm 5c$. In terms of angular expansion rate, the maximum values of our survey are limited to roughly $<3000 \mu\text{as yr}^{-1}$ due to our limited angular resolution and observing cadence. The fastest measured angular speed in our survey to date is $2941 \pm 109 \mu\text{as yr}^{-1}$ in the jet of the nearby ($z = 0.033$) AGN 3C 120 (Lister et al. 2013).

We note that there are two accelerating features whose apparent speeds exceeded $40c$ at some point during our monitoring (the speeds listed in Table 5 are the instantaneous speeds at the midpoint epoch). These are feature id = 3 in PKS 0805–07, which had an initial speed of $\sim 50c$ in 1996 and subsequently decelerated to $\sim 20c$ and feature id = 16 in PKS 1510–08, which attained a maximum speed of $\sim 42c$ before fading in 2011.

In Lister et al. (2013), we identified 38 jet features as having slow pattern speeds, i.e., they had non-accelerating, angular speeds smaller than $20 \mu\text{as yr}^{-1}$ and a speed that was at least 10 times slower than the fastest feature in the same jet. With the addition of our new kinematics analysis, we now have identified a total of 63 such features (5% of all robust features), which we have flagged in Table 4.

3.2. Jet Speed and Gamma-ray Emission

The most recent catalog of gamma-ray sources detected by the *Fermi*-LAT instrument (3FGL: Acero et al. 2015) is highly dominated by blazars, due to the strong influence of relativistic beaming on jet gamma-ray emission. As discussed by Lister et al. (2015), based on data from Lister et al. (2013), nearly all of the fastest AGN jets in the MOJAVE program have been detected by *Fermi*, indicating a strong correlation between jet speed and the gamma-ray Doppler boosting factor. Several fast jets have yet to be detected by *Fermi* because their spectral energy distribution peaks below the 100 MeV threshold of the LAT. In Figure 7, we show the distributions of fastest and median jet speed for our full MOJAVE sample (top panels), and our 1.5 Jy flux density-limited sample (lower panels), including our new speed data. In terms of the fastest measured speeds, there is a less than 0.02% probability that the LAT and

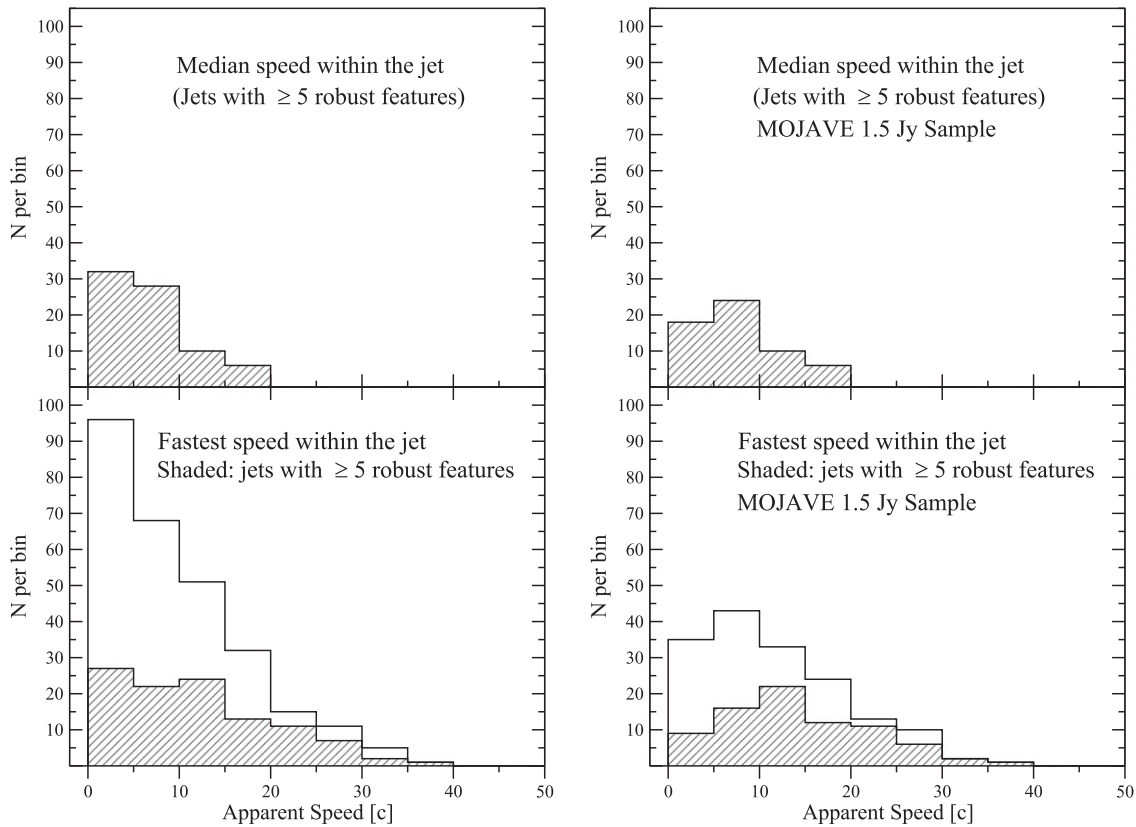


Figure 6. Distributions of median and fastest speeds for the jets in our sample with measured speeds and redshifts. The shaded histograms include only those jets with at least 5 robust features. The left-hand panels contain all speeds for all AGNs measured by the MOJAVE program to date, while the right-hand panels are for AGNs in the complete flux density-limited 1.5 Jy MOJAVE sample only.

non-LAT sub-samples come from the same parent distribution according to Kolmogorov–Smirnov tests, in the case of both the full and MOJAVE 1.5 Jy flux density-limited samples.

3.3. Apparent Jet Speed Versus Redshift

In Figure 8, we plot maximum measured apparent jet speed versus redshift for all jets measured by the MOJAVE program to date. For completeness, we also plot as lower limits equal to $0.015c$ those AGN jets for which we were not able to identify any robust features for kinematics analysis, despite having at least five VLBA epochs.

This collection of AGNs represents those selected from the MOJAVE radio flux density-limited, gamma-ray-selected, and low-luminosity blazar samples. As such, it spans a very large range of redshift ($0.004 < z < 3.4$) and apparent radio luminosity (10^{23} – 10^{29} W Hz^{-1}). At low redshifts, lower-luminosity radio galaxies and BL Lac objects dominate, since high luminosity quasars are much rarer in the general AGN population (see, e.g., discussion of luminosity functions in Urry & Padovani 1995).

Our monthly observing cadence, total time baseline, and angular resolution set limits on the maximum and minimum observable angular proper motions, which we plot as dotted lines in Figure 8. We note that for $z > 0.3$ our survey is sensitive to extremely high jet speeds, yet there is a distinct upper envelope to the distribution. Since there must be a portion of the ultrafast AGN jet population that is oriented close to our line of sight, and such jets will be extremely bright due to enormous Doppler boosting of their radio flux density,

the upper envelope in Figure 8 is a clear indication of an upper limit of ~ 50 in the bulk Lorentz factors in the parsec-scale radio-emitting regions of AGN jets.

In Figure 8, we have also indicated with filled symbols those AGNs that have been detected at TeV gamma-ray energies. To date, there are roughly 60 known AGNs in this category¹², and one-third of these have been monitored by the MOJAVE program. Our plot illustrates the two types of known TeV AGNs: (1) BL Lac objects that have high TeV fluxes due to their proximity and the fact that their spectral energy distributions are peaked at high frequencies and (2) other BL Lac objects and quasars with very fast jet speeds, indicating unusually high amounts of relativistic beaming. To date, there have been no TeV detections of AGNs past $z \simeq 1$ due to the strong attenuation of high-energy gamma-rays by pair production off of the extragalactic background light. This would suggest that a promising region to search for new TeV-emitting AGNs is in the top left corner of Figure 8, where $\beta_{\text{app}} > 10c$ and $z < 0.35$. This region includes the BL Lac objects PKS 0754+100, OJ 287 and OJ 049, the quasars 3C 273, OI 061, 4C +49.22, TXS 1700+685, and PKS 1725+044, the AGNs 3C 111 and 3C 120, and the NLSY1 1H 0323+342.

We have measured maximum jet speeds for nine high spectral peaked (HSP) BL Lacs to date in the MOJAVE program (triangle symbols in Figure 8), all but one of which (7C 1055+5644) have been detected at TeV energies. Their apparent speeds are sub-luminal, ranging from $0.03c \pm 0.01c$ in ON 325 to $2.6c \pm 1.1c$ in OQ 240. This is in agreement with

¹² <http://tevcat.uchicago.edu>

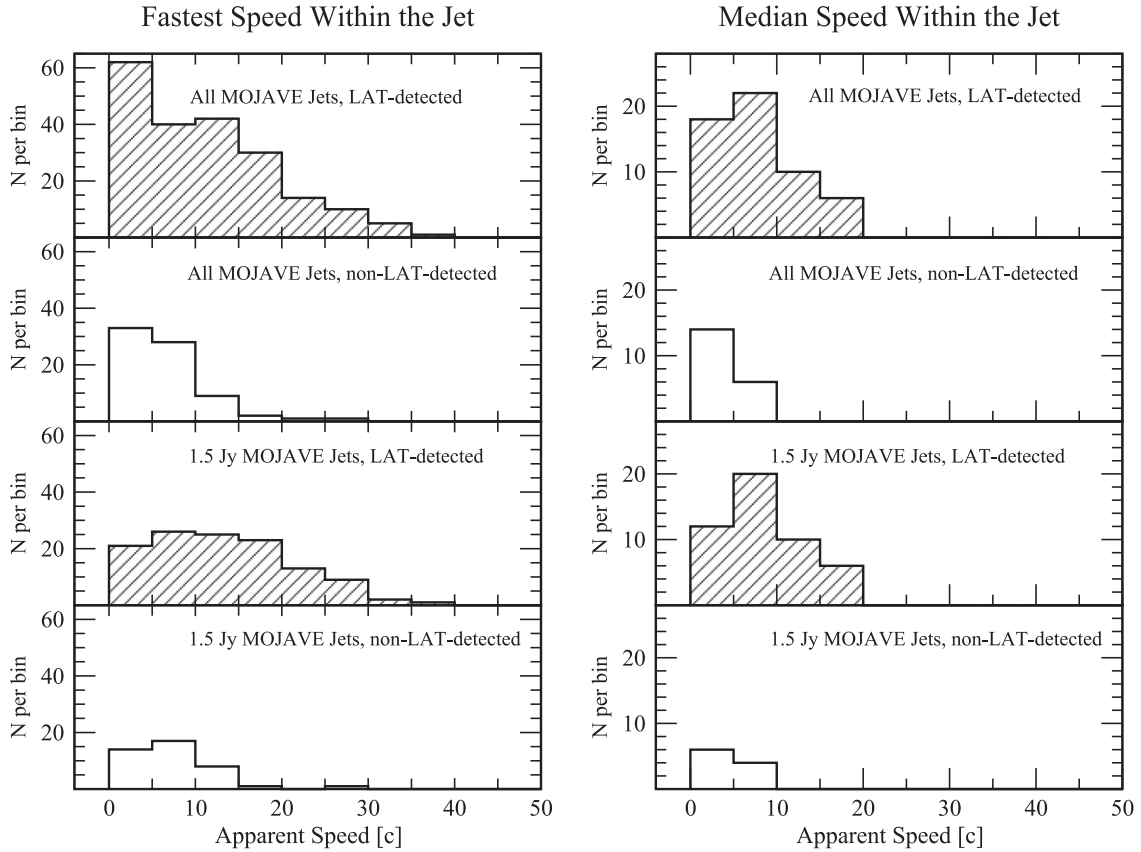


Figure 7. Jet speed distributions for various MOJAVE sub-samples. The left-hand panels show the distributions for the fastest measured jet speed, while the right-hand panels show the distributions of median jet speed for jets having at least 5 robust features. The topmost two panels show all speeds measured by the MOJAVE program to date for *Fermi* LAT detected (shaded) and non-LAT detected AGN (unshaded), respectively. The bottom two panels show the speeds of the LAT detected (shaded) and non-LAT detected (unshaded) AGN, respectively, in the flux density-limited 1.5 Jy MOJAVE sample.

previous multi-epoch VLBI observations of a small number of TeV blazar jets at 8 GHz by Piner et al. (2008). We are currently obtaining multi-epoch MOJAVE observations of a much larger set of HSP AGN detected by the *Fermi* γ -ray observatory in order to further investigate their speed distribution.

3.4. Radio-loud Narrow-lined Seyfert I AGNs

The third *Fermi* AGN catalog (Ackermann et al. 2015) is comprised primarily of blazars, reflecting the strong influence of relativistic boosting on AGN gamma-ray emission. Apart from a small handful of nearby radio galaxies, the only other major AGN category in the *Fermi* catalog is that of radio-loud NLSY1 galaxies. To date, fewer than a dozen gamma-ray-loud NLSY1s have been identified (Berton et al. 2015), but they are interesting objects since they appear to have many of the same characteristics as blazars, despite the fact that Seyferts are typically found in spiral host galaxies. The orientation of NLSY1s renders their broad line region visible, yet they display narrow permitted lines ($<2000 \text{ km s}^{-1}$) due to their small black hole masses ($10^6\text{--}10^8 M_\odot$). Therefore, in order to produce blazar-like gamma-ray and radio luminosities, they must have high Eddington ratios. This has led to speculation that they might be young radio jets that have recently been fueled and are perhaps the low-black-hole mass, low viewing angle versions of compact symmetric AGNs (e.g., Kawakatu et al. 2009).

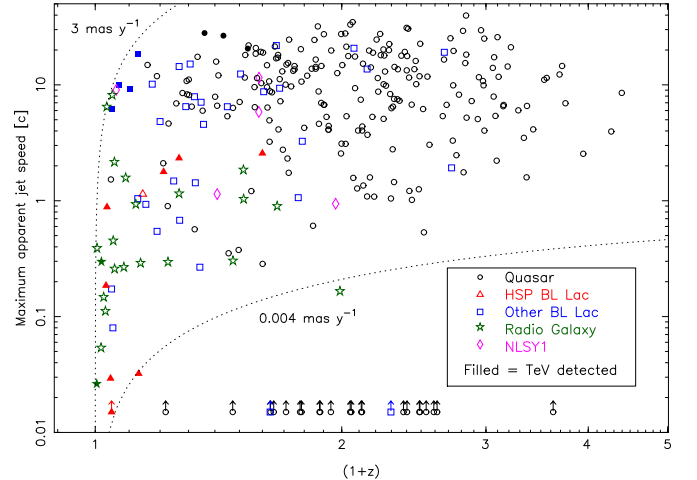


Figure 8. Log-log plot of maximum apparent jet speed vs. redshift for all jets measured by the MOJAVE program to date. The dotted lines delimit the approximate survey limits on angular proper motions, based on our angular resolution, temporal coverage, and observing cadence. Black circles correspond to quasars, red triangles to high spectral peaked BL Lac objects, blue squares to other BL Lacs, green stars to radio galaxies, and purple diamonds to NLSY1s. Filled symbols indicate AGNs that have been detected at TeV gamma-ray energies. Jets for which we were not able to identify any robust features for kinematic analysis are plotted as lower limits with an arbitrary value of $0.015c$.

We have analyzed the parsec-scale jet kinematics of five NLSY1s in the MOJAVE sample, all of which have been detected in gamma-rays by *Fermi*. Other groups (D’Ammando

et al. 2012, 2013) have published simple one-dimensional radial-fit speeds for two of these AGNs using our MOJAVE data. We find that three of the NLSY1s show highly superluminal jet features ($9.0c \pm 0.3c$ for 1H 0323+342, $5.8c \pm 0.9c$ for SBS 0846+513, and $11.5c \pm 1.5c$ for PMN J0948+0022). The other two show kinematics consistent with sub-luminal speeds: $0.9c \pm 0.3c$ for 4C +04.42 and $1.1c \pm 0.4c$ for 1502+036. With this small sample size, it is impossible to draw any general conclusions about the population, but the high speeds indicate that at least some radio-loud NLSY1s have Lorentz factors in excess of 10, and viewing angles less than 10° consistent with those of BL Lac objects and flat-spectrum radio quasars.

4. SUMMARY

We have presented 1625 new 15 GHz VLBA contour images of 295 AGNs from the MOJAVE and 2 cm VLBA surveys, spanning epochs from 1994 August 31 to 2013 August 20. For the 274 AGNs with at least five temporally spaced epochs, we have analyzed the kinematics of individual bright features in the jets, and produced time-lapse linearly interpolated movies that show their parsec-scale evolution. At least half of the well-sampled features show non-radial and/or accelerating trajectories, indicating that non-ballistic motion is common. Since it is impossible to extrapolate any accelerations that occurred before our monitoring period or below our resolution level, we are only able to determine reliable ejection dates for $\sim 24\%$ of those features with significant proper motions.

The distribution of maximum jet speeds in all 295 AGNs measured by our program to date is peaked below $5c$ with very few jets with apparent speeds above $30c$. The fastest instantaneous measured speed in our survey is $\sim 50c$ measured in the jet of the quasar PKS 0805–07. The form of the speed distribution is indicative of a moderately steep power-law distribution of Lorentz factors in the parent population, ranging up to $\Gamma \simeq 50$. An upper envelope in the maximum speed versus redshift distribution provides additional evidence of this upper limit to AGN jet speeds in the parsec-scale regime.

The *Fermi*-LAT-detected gamma-ray AGNs in our sample have, on average, higher jet speeds than non-LAT-detected AGNs, indicating a strong correlation between parsec-scale jet speed and the gamma-ray Doppler boosting factor. The known TeV gamma-ray emitting AGNs in our sample are clustered in two regions: namely, at low redshift/low speeds, and at moderate redshift/high speeds. We have identified 11 AGNs in the latter region ($\beta_{\text{app}} > 10c$, $z < 0.35$) that are strong candidates for future TeV gamma-ray detection. Our AGN sample also includes five gamma-ray-loud NLSY1s, a rare type of low-black hole mass, high-accretion rate AGN that has been identified as a distinct population of gamma-ray emitting AGNs by *Fermi*. Of the five NLSY1s, three show highly superluminal jet motions, while the others have sub-luminal speeds. This indicates that radio-loud NLSY1 AGNs can possess powerful jets with high Lorentz factors and low viewing angles that are consistent with typical radio-loud BL Lac objects and flat-spectrum radio quasars.

The MOJAVE project was supported by NASA-*Fermi* GI grants NNX08AV67G, NNX12A087G, and NNX15AU76G. The National Radio Astronomy Observatory is a facility of the National Science Foundation operated under cooperative agreement by Associated Universities, Inc. E.R. acknowledges

partial support by the Spanish MINECO project AYA2012-38491-C02-01 and by the Generalitat Valenciana project PROMETEOII/2014/057. M.F.A. was supported in part by NASA-*Fermi* GI grants NNX09AU16G, NNX10AP16G, NNX11AO13G, and NNX13AP18G and NSF grant AST-0607523. Y.Y.K. and A.B.P. were supported by the Russian Foundation for Basic Research (project 13-02-12103) and the Basic Research Program P-7 of the Presidium of the Russian Academy of Sciences. T.S. was supported by the Academy of Finland projects 274477 and 284495. This work made use of the Swinburne University of Technology software correlator (Deller et al. 2011), developed as part of the Australian Major National Research Facilities Programme and operated under licence.

APPENDIX NOTES ON INDIVIDUAL AGNS

0006 + 061 (CRATES J0009+0628): Rau et al. (2012) obtained $z < 1.35$ for this BL Lac based on photometric redshift estimates. It is one of two possible 3LAC associations for 3FGL J0009.1+0630 listed by Ackermann et al. (2015). The other is CRATES J0009+0625.

0048–097 (PKS 0048–09): the jet structure is too compact at 15 GHz to reliably measure any robust features.

0106 + 678 (4C +67.04): this BL Lac has an inward-moving feature (id = 4) located 1.3 mas from the core.

0109 + 224 (S2 0109+22): The redshift $z = 0.265$ from Shaw et al. (2012) that we have used to derive the apparent speeds is claimed to be incorrect according to Paiano et al. (2016), who find a lower limit of $z > 0.35$, based on new optical spectroscopy.

0108 + 388 (CGRABS J0111+3906): multi-frequency VLBA images of this radio galaxy indicate a two-sided jet morphology; however, the precise location of the core is not known (Marr et al. 2001). We therefore used the southwesternmost bright feature in the images as a reference point for the kinematics analysis.

0111 + 021 (UGC 00773): with the addition of one more epoch in 2013 since our Lister et al. (2013) analysis, we now classify a total of four features in this BL Lac jet as inward-moving.

0118–272 (OC –230.4): this BL Lac has a very slow inward-moving feature ($25 \pm 7 \mu\text{s yr}^{-1}$) located 0.5 mas from the core. Shaw et al. (2013b) found $z < 0.558$ based on intergalactic absorption features in the optical spectrum.

0119 + 041 (PKS 0119+041): the jet structure is too compact at 15 GHz to reliably measure any robust features.

0141 + 268 (TXS 0141+268): Shaw et al. (2009) reported $z < 2.41$ for this BL Lac object.

0219 + 428 (3C 66A): the two innermost jet features of this BL Lac object have robust inward motions. No reliable spectroscopic redshift has been published to date. Stadnik & Romani (2014) reported $z > 0.42$ based on the host galaxy magnitude.

0235 + 164 (AO 0235+164): the jet structure is too compact at 15 GHz to reliably measure any robust features.

0300 + 470 (4C +47.08): no reliable spectroscopic redshift exists for this BL Lac (Shaw et al. 2013b), which contains an inward-moving feature (id = 1) at 0.8 mas from the core feature.

0301–243 (PKS 0301–243): the innermost feature of this BL Lac jet has a robust inward motion.

0321 + 340: a detailed study on this AGN, including single-dish radio flux density data is presented by Karamanavis (2015) and Fuhrmann et al. (2016).

0346 + 800 (S5 0346+80): no reliable spectroscopic redshift exists in the literature for this BL Lac object.

0355 + 508: The core location in the last epoch (2012 January 14) is uncertain due to the appearance of a bright polarized feature.

0414–189 (PKS 0414–189): the jet structure is too compact at 15 GHz to reliably measure any robust features.

0440–003 (NRAO 190): this quasar is one of two possible 3LAC associations for 3FGL J0442.6–0017 listed by Ackermann et al. (2015). The other is a radio-quiet X-ray source (1RXS J044229.8–001823), which is 2.3 arcmin from the LAT position (as compared to 0.21 arcmin for 0440–003).

0454 + 844 (S5 0454+84): the value of $z = 1.340$ from Rector & Stocke (2001) is a lower limit, and the value $z = 0.113$ from Lawrence et al. (1996) is considered tentative.

0506 + 056 (TXS 0506+056): as described by Lister et al. (2013), this BL Lac object has an inward-moving feature (id = 3) at 1.3 mas from the core feature.

0615 + 820 (S5 0615+82): the jet structure is too compact to reliably identify the core position at all epochs, or to classify any jet features as robust.

0640 + 090 (PMN J0643+0857): the position of the core feature in this low galactic latitude quasar is uncertain, and we considered none of the jet features to be robust. We used the highest brightness temperature feature as a reference point for the kinematics analysis.

0646 + 600 (S4 0646+60): this is a high-frequency peaked quasar (Dallacasa et al. 2000) with two-sided parsec-scale radio structure, meeting the criteria for a compact symmetric object. The outermost jet feature at ~ 3 mas from the core has slow but significant inward motion ($10 \pm 0.7 \mu\text{as yr}^{-1}$).

0727–115 (PKS 0727–11): the radio morphology of this quasar was too compact and complex to identify any suitably robust jet features for kinematic analysis.

0742 + 103 (J0745+1011): this high-redshift quasar has a gigahertz-peaked spectrum and jet emission to both the NW and SE of the brightest feature in the maps. The core location is unknown, making it impossible to reliably cross-register the maps across the epochs or to classify any jet features as robust.

0743–006 (J0745-0044): the position of the core in this gigahertz-peaked spectrum AGN is uncertain. We used the position of the most compact feature in the maps as a reference point for the kinematics analysis.

0745 + 241 (S3 0745+24): the jet feature at 3.4 mas from the core (id = 11) has significant inward motion.

0804 + 499 (OJ 508): this quasar is one of two possible 3LAC associations for 3FGL J0807.9+4946 listed by Ackermann et al. (2015). The other is a double-lobed radio galaxy (J0807+4946) with no visible radio core in the FIRST survey image, located 7.74 arcmin from the LAT position (as compared to 0.77 arcmin for OJ 508).

0814 + 425 (OJ 425): the value of $z = 0.53$ from Linford et al. (2012) is uncertain, while Shaw et al. (2013b) reported $z < 2.47$ and Sbarufatti et al. (2005a) obtained $z > 0.75$. As described in Lister et al. (2013), the innermost feature (id = 5) has robust inward motion.

0821 + 394 (4C +39.23): this quasar is one of two possible 3LAC associations for 3FGL J0824.9+3916 listed by

Ackermann et al. (2015). The other is the compact steep-spectrum quasar 4C +39.23B (Dallacasa et al. 2002).

0823–223 (PKS 0823–223): the value of $z = 0.91$ from Falomo (1990) is a lower limit. We could not identify any suitably robust jet features for kinematic analysis in this BL Lac.

0946 + 006 (PMN J0948+0022): using additional VLBA epochs on this narrow-lined Seyfert I galaxy obtained since our Lister et al. (2013) analysis, we now classify two jet features as robust.

0954 + 658 (S4 0954+65): the widely cited redshift of $z = 0.367$ by Stickel et al. (1993a) has been drawn in to question by Landoni et al. (2015), who find a featureless spectrum and a limit $z > 0.45$ based on non-detection of the host galaxy. We note that for $z > 1.7$, our apparent speed measurement of $\mu = 672 \pm 50 \mu\text{as yr}^{-1}$ would yield $v_{\text{app}} \simeq 50c$, a value larger than the fastest measured speed of any AGN in the MOJAVE program.

1030 + 415 (S4 1030+41): the jet structure is too compact at 15 GHz to reliably measure any robust features.

1101 + 384 (Mrk 421): the innermost feature (id = 8) in this nearby BL Lac jet has slow but significant inward motion.

1124–186 (PKS 1124–186): the jet structure is too compact at 15 GHz to reliably measure any robust features.

1128 + 385 (B2 1128+38): the outermost feature in this jet has significant inward motion.

1144 + 402 (S4 1144+40): this quasar is one of two possible associations listed for 3FGL J1146.8+3958 by Ackermann et al. (2015). The other is the weak radio source NVSS J114653+395751, which showed less than 1 mJy of correlated VLBA flux density at 1.4 GHz in the mJIVE-20 survey of Deller & Middelberg (2014).

1148–001 (4C–00.47): the position of the core in this high-redshift ($z = 1.98$) quasar is uncertain. We used the position of the most compact feature in the maps as a reference point for the kinematics analysis.

1213–172 (PKS 1213–17): this AGN is located very close on the sky to a bright star and has no reliable spectroscopic redshift in the literature.

1219 + 044 (4C +04.42): Yao et al. (2015) have classified this AGN as a NLSY1 galaxy, based on its optical spectrum from the SDSS-BOSS survey.

1219 + 285 (W Comae): a feature (id = 17) located ~ 1 mas from the core has significant inward motion.

1243–072 (PKS 1243–072): this quasar is listed as an association in the third EGRET catalog (Hartman et al. 1999), but not in any of the *Fermi* catalogs.

1324 + 224 (B2 1324+22): the jet structure is too compact at 15 GHz to reliably measure any robust features.

1329–126 (PMN J1332–1256): we did not identify any suitably robust jet features for kinematic analysis in this quasar.

1331 + 170 (OP 151): this quasar is listed as an association in the third EGRET catalog (Hartman et al. 1999), but was not confirmed in the revised EGRET catalog of Casandjian & Grenier (2008). It is not listed in any of the *Fermi* catalogs.

1413 + 135 (J1415+1320): this unusual BL Lac object resides in a spiral host galaxy (Perlman et al. 2002), and has a two-sided jet structure (Gugliucci et al. 2005). The core is associated with the brightest feature in our VLBA images.

1435 + 638 (VIPS 0792): we presumed the radio core to lie at the northernmost end of the jet, based on the VIPS 5 GHz VLBA map of Helmboldt et al. (2007).

1458 + 718 (3C 309.1): this compact steep-spectrum quasar has a group of inward-moving jet features located 23 mas south of the core. On larger scales, the radio morphology suggests a twisting (helical) jet structure (Ludke et al. 1998).

1508–055 (PKS 1508–05): this compact steep-spectrum quasar is listed as a gamma-ray association in the 3LAC catalog (Ackermann et al. 2015). We find no significant proper motion in two robust jet features.

1509 + 054 (PMN J1511+0518): this gigahertz-peaked spectrum radio galaxy has a two-sided jet morphology (Hovatta et al. 2014), and a statistically significant inward-moving feature (id = 2). However, it is possible that we are merely seeing changes in the internal brightness distribution of unresolved features in this jet, and not true apparent motion.

1510–089 (PKS 1510–08): many of the features in this jet (e.g., id = 8, 15, 16) show complex kinematics, including a kink in their trajectories at 2 mas from the core that are not well approximated by a constant acceleration model.

1514 + 004 (PKS 1514+00): the brightness profile along this nearby ($z = 0.052$) radio galaxy jet is very smooth, and not well-represented by discrete Gaussian features.

1519–273 (PKS 1519–273): the jet structure is too compact at 15 GHz to reliably measure any robust features.

1529–131 (PMN J1532–1319): there is no known optical counterpart for this AGN. We could not identify any suitably robust jet features for kinematic analysis.

1548 + 056 (4C +05.64): the southernmost jet feature, which we presume marks the location of the core, faded in the 2005 October 29 epoch, thus we have not used this epoch in our kinematics analysis.

1611 + 343 (DA 406): the morphology and complex kinematics of this quasar jet suggest that it is bending into our line of sight several milliarcseconds from the core, and that we may be seeing features from both the front and back sides of the jet.

1622–253 (PKS 1622–253): the jet structure is too compact at 15 GHz to reliably measure any robust features.

1637 + 826 (NGC 6251): in Lister et al. (2013), we reported an inward-moving feature (id = 8) in this radio galaxy jet. With the addition of two new epochs, this feature no longer shows any statistically significant motion.

1656 + 053 (PKS 1656+053): we could not identify any suitably robust jet features for kinematic analysis in this quasar.

1656 + 477 (S4 1656+47): we could not identify any suitably robust jet features for kinematic analysis in this quasar.

1739 + 522 (4C +51.37): the jet structure is too compact at 15 GHz to reliably measure any robust features.

1741–038 (PKS 1741–03): the jet structure is too compact at 15 GHz to reliably measure any robust features.

1749 + 701 (S4 1749+70): all of the jet features in this BL Lac show transverse motions, including the innermost feature, which evolves from a position angle of -50° to -75° over a 15-year period.

1842 + 681 (S4 1842+68): we could not identify any suitably robust jet features for kinematic analysis in this quasar.

1923 + 210 (PKS B1923+210): this low galactic latitude AGN has a featureless optical spectrum, as reported by Titov et al. (2011).

1958–179 (PKS 1958–179): the jet structure is too compact at 15 GHz to reliably measure any robust features.

2021 + 614 (OW 637): the core feature location in this jet remains uncertain (Lister et al. 2009b). The two outermost

features (ID = 1 and 2) have inward motions, albeit with very slow speeds (8.8 and $19.6 \mu\text{as yr}^{-1}$, respectively).

2023 + 335 (B2 2023+33): this low galactic latitude quasar is strongly affected by interstellar scattering (Pushkarev et al. 2013). The radio structure was too complex to reliably cross-identify any robust features for kinematic analysis. Examples of this kind of strong scattering are extremely rare among bright, compact AGNs (Pushkarev & Kovalev 2015).

2047 + 098 (PKS 2047+098): this AGN has no known optical counterpart. The value $z = 0.01513$ listed in the 2LAC (Ackermann et al. 2011a) data table is unreferenced. Its jet has two inward-moving features (id = 2 and 3).

2128–123 (PKS 2128–12): the overall kinematics and polarization structure of this quasar underwent sudden and dramatic changes in 2009.

2144 + 092 (PKS 2144+092): we could not identify any suitably robust jet features for kinematic analysis in this quasar.

2155 + 312 (B2 2155+31): we could not identify any suitably robust jet features for kinematic analysis in this quasar.

2230 + 114 (CTA 102): by using additional VLBA epochs obtained since Lister et al. (2013), we now identify two robust inward-moving features (id = 1 and 4) in this quasar jet. These were also identified in the multi-frequency VLBI kinematic analysis of this AGN by Fromm et al. (2013).

2234 + 282 (CTD 135): An et al. (2016) have claimed this BL Lac to be a compact symmetric object with a two-sided jet, based on a comparison of VLBA maps at 8.4 and 15 GHz. However, the angular resolution is insufficient to determine a spectral index for the putative core feature. For the purposes of our kinematic analysis, we have identified the core with the northernmost jet feature in our maps.

2247–283 (PMN J2250–2806): by using additional VLBA epochs we have obtained since Lister et al. (2013), we are able to identify one robust feature in this quasar jet.

2356 + 196 (PKS 2356+196): this quasar is listed as an association in the third EGRET catalog (Hartman et al. 1999), but was not confirmed in the revised EGRET catalog of Casandjian & Grenier (2008). It is not listed in any of the *Fermi* catalogs.

REFERENCES

- Abazajian, K., Adelman-McCarthy, J. K., Agüeros, M. A., et al. 2004, *AJ*, **128**, 502
- Abazajian, K., Adelman-McCarthy, J. K., Agüeros, M. A., et al. 2005, *AJ*, **129**, 1755
- Acero, F., Ackermann, M., Ajello, M., et al. 2015, *ApJS*, **218**, 23
- Ackermann, M., Ajello, M., Allafort, A., et al. 2011a, *ApJ*, **743**, 171
- Ackermann, M., Ajello, M., Allafort, A., et al. 2011b, *ApJ*, **743**, 171
- Ackermann, M., Ajello, M., Atwood, W. B., et al. 2015, *ApJ*, **810**, 14
- Adelman-McCarthy, J. K., Agüeros, M. A., Allam, S. S., et al. 2006, *ApJS*, **162**, 38
- Adelman-McCarthy, J. K., Agüeros, M. A., Allam, S. S., et al. 2008, *ApJS*, **175**, 297
- Agudo, I., Bach, U., Krichbaum, T. P., et al. 2007, *A&A*, **476**, L17
- Akiyama, M., Ueda, Y., Ohta, K., Takahashi, T., & Yamada, T. 2003, *ApJS*, **148**, 275
- Albert, J., Aliu, E., Anderhub, H., et al. 2007, *ApJL*, **667**, L21
- Aldcroft, T. L., Bechtold, J., & Elvis, M. 1994, *ApJS*, **93**, 1
- An, T., Cui, Y.-Z., Gabányi, K. É., et al. 2016, *AN*, **337**, 65
- Antonucci, R. R. J., Hickson, P., Miller, J. S., & Olszewski, E. W. 1987, *AJ*, **93**, 785
- Baker, J. C., Hunstead, R. W., Kapahi, V. K., & Subrahmanya, C. R. 1999, *ApJS*, **122**, 29
- Baldwin, J. A., Wampler, E. J., & Burbidge, E. M. 1981, *ApJ*, **243**, 76
- Bartel, N., Shapiro, I. I., Huchra, J. P., & Kuhr, H. 1984, *ApJ*, **279**, 112

- Bauer, F. E., Condon, J. J., Thuan, T. X., & Broderick, J. J. 2000, *ApJS*, **129**, 547
- Berton, M., Foschini, L., Ciroi, S., et al. 2015, arXiv:1512.00171
- Best, P. N., Peacock, J. A., Brookes, M. H., et al. 2003, *MNRAS*, **346**, 1021
- Blandford, R. D., & Königl, A. 1979, *ApJ*, **232**, 34
- Blandford, R. D., McKee, C. F., & Rees, M. J. 1977, *Natur*, **267**, 211
- Boisse, P., & Bergeron, J. 1988, *A&A*, **192**, 1
- Boksenberg, A., Briggs, S. A., Carswell, R. F., Schmidt, M., & Walsh, D. 1976, *MNRAS*, **177**, 43
- Browne, I. W. A., Savage, A., & Bolton, J. G. 1975, *MNRAS*, **173**, 87
- Burbidge, E. M. 1970, *ApJL*, **160**, L33
- Carangelo, N., Falomo, R., Kotilainen, J., Treves, A., & Ulrich, M.-H. 2003, *A&A*, **412**, 651
- Carpenter, B., Ojha, R., D'Ammando, F., Orienti, M., & Cheung, C. C. 2014, *ATel*, **5838**, 1
- Casandjian, J.-M., & Grenier, I. A. 2008, *A&A*, **489**, 849
- Chavushyan, V., Mujica, R., Gorshkov, A. G., et al. 2001, *ARep*, **45**, 79
- Cohen, M. H., Meier, D. L., Arshakian, T. G., et al. 2015, *ApJ*, **803**, 3
- Cohen, R. D., Smith, H. E., Junkkarinen, V. T., & Burbidge, E. M. 1987, *ApJ*, **318**, 577
- Dallacasa, D., Stanghellini, C., Centonza, M., & Fanti, R. 2000, *A&A*, **363**, 887
- Dallacasa, D., Tinti, S., Fanti, C., et al. 2002, *A&A*, **389**, 115
- D'Ammando, F., Orienti, M., Finke, J., et al. 2012, *MNRAS*, **426**, 317
- D'Ammando, F., Orienti, M., Doi, A., et al. 2013, *MNRAS*, **433**, 952
- de Grijs, M. H. K., Keel, W. C., Miley, G. K., Goudfrooij, P., & Lub, J. 1992, *A&AS*, **96**, 389
- Deller, A. T., & Middelberg, E. 2014, *AJ*, **147**, 14
- Deller, A. T., Briske, W. F., Phillips, C. J., et al. 2011, *PASP*, **123**, 275
- Denicolò, G., Terlevich, R., Terlevich, E., et al. 2005, *MNRAS*, **356**, 1440
- di Serego-Alighieri, S., Danziger, I. J., Morganti, R., & Tadhunter, C. N. 1994, *MNRAS*, **269**, 998
- Drinkwater, M. J., Webster, R. L., Francis, P. J., et al. 1997, *MNRAS*, **284**, 85
- Dunlop, J. S., Peacock, J. A., Savage, A., et al. 1989, *MNRAS*, **238**, 1171
- Eckart, A., Witzel, A., Biermann, P., et al. 1986, *A&A*, **168**, 17
- Ellison, S. L., Yan, L., Hook, I. M., et al. 2001, *A&A*, **379**, 393
- Eracleous, M., & Halpern, J. P. 1994, *ApJS*, **90**, 1
- Eracleous, M., & Halpern, J. P. 2004, *ApJS*, **150**, 181
- Falomo, R. 1990, *ApJ*, **353**, 114
- Falomo, R., Scarpa, R., & Bersanelli, M. 1994, *ApJS*, **93**, 125
- Fomalont, E. B. 1999, in ASP Conf. Ser. 180, Synthesis Imaging in Radio Astronomy II, ed. G. B. Taylor, C. L. Carilli, & R. A. Perley (San Francisco, CA: ASP), 301
- Fomalont, E. B., Frey, S., Paragi, Z., et al. 2000, *ApJS*, **131**, 95
- Fricke, K. J., Kollatschny, W., & Witzel, A. 1983, *A&A*, **117**, 60
- Fromm, C. M., Ros, E., Peruchio, M., et al. 2013, *A&A*, **551**, A32
- Fuhrmann, L., Karamanavis, V., Komossa, S., et al. 2016, *MNRAS*, submitted
- Furniss, A., Fumagalli, M., Danforth, C., Williams, D. A., & Prochaska, J. X. 2013, *ApJ*, **766**, 35
- Gelderman, R., & Whittle, M. 1994, *ApJS*, **91**, 491
- Gugliucci, N. E., Taylor, G. B., Peck, A. B., & Giroletti, M. 2005, *ApJ*, **622**, 136
- Halpern, J. P., & Eracleous, M. 1997, *IAUC*, **6639**, 2
- Halpern, J. P., Eracleous, M., & Mattox, J. R. 2003, *AJ*, **125**, 572
- Hartman, R. C., Bertsch, D. L., Bloom, S. D., et al. 1999, *ApJS*, **123**, 79
- Healey, S. E., Romani, R. W., Cotter, G., et al. 2008, *ApJS*, **175**, 97
- Heidt, J., Tröller, M., Nilsson, K., et al. 2004, *A&A*, **418**, 813
- Helmboldt, J. F., Taylor, G. B., Tremblay, S., et al. 2007, *ApJ*, **658**, 203
- Henstock, D. R., Browne, I. W. A., Wilkinson, P. N., & McMahon, R. G. 1997, *MNRAS*, **290**, 380
- Hewett, P. C., Foltz, C. B., & Chaffee, F. H. 1995, *AJ*, **109**, 1498
- Hewett, P. C., & Wild, V. 2010, *MNRAS*, **405**, 2302
- Ho, L. C., & Kim, M. 2009, *ApJS*, **184**, 398
- Homan, D. C., Kadler, M., Kellermann, K. I., et al. 2009, *ApJ*, **706**, 1253
- Homan, D. C., Lister, M. L., Kellermann, K. I., et al. 2003, *ApJL*, **589**, L9
- Homan, D. C., Lister, M. L., Kovalev, Y. Y., et al. 2015, *ApJ*, **798**, 134
- Homan, D. C., Ojha, R., Wardle, J. F. C., et al. 2002, *ApJ*, **568**, 99
- Hook, I. M., McMahon, R. G., Irwin, M. J., & Hazard, C. 1996, *MNRAS*, **282**, 1274
- Hovatta, T., Lister, M. L., Aller, M. F., et al. 2012, *AJ*, **144**, 105
- Hovatta, T., Aller, M. F., Aller, H. D., et al. 2014, *AJ*, **147**, 143
- Huchra, J. P., Vogeley, M. S., & Geller, M. J. 1999, *ApJS*, **121**, 287
- Hunstead, R. W., Murdoch, H. S., & Shobbrook, R. R. 1978, *MNRAS*, **185**, 149
- Hunter, S. D., Bertsch, D. L., Dingus, B. L., et al. 1993, *ApJ*, **409**, 134
- Jackson, N., & Browne, I. W. A. 1991, *MNRAS*, **250**, 414
- Jauncey, D. L., Batty, M. J., Wright, A. E., Peterson, B. A., & Savage, A. 1984, *ApJ*, **286**, 498
- Jones, D. H., Saunders, W., Read, M., & Colless, M. 2005, *PASA*, **22**, 277
- Jones, D. H., Read, M. A., Saunders, W., et al. 2009, *MNRAS*, **399**, 683
- Junkkarinen, V. 1984, *PASP*, **96**, 539
- Karamanavis, V. 2015, PhD thesis, Max-Planck-Institut für Radioastronomie
- Kawakatu, N., Nagao, T., & Woo, J.-H. 2009, *ApJ*, **693**, 1686
- Kellermann, K. I., Vermeulen, R. C., Zensus, J. A., & Cohen, M. H. 1998, *AJ*, **115**, 1295
- Kellermann, K. I., Lister, M. L., Homan, D. C., et al. 2004, *ApJ*, **609**, 539
- Komatsu, E., Dunkley, J., Nolte, M. R., et al. 2009, *ApJS*, **180**, 330
- Labiano, A., Barthel, P. D., O'Dea, C. P., et al. 2007, *A&A*, **463**, 97
- Landoni, M., Falomo, R., Treves, A., et al. 2012, *A&A*, **543**, A116
- Landoni, M., Falomo, R., Treves, A., Scarpa, R., & Reverte Payá, D. 2015, *AJ*, **150**, 181
- Lanzetta, K. M., Wolfe, A. M., Turnshek, D. A., et al. 1991, *ApJS*, **77**, 1
- Lawrence, C. R., Pearson, T. J., Readhead, A. C. S., & Unwin, S. C. 1986, *AJ*, **91**, 494
- Lawrence, C. R., Zucker, J. R., Readhead, A. C. S., et al. 1996, *ApJS*, **107**, 541
- Linford, J. D., Taylor, G. B., Romani, R. W., et al. 2012, *ApJ*, **744**, 177
- Lister, M. L., Aller, M. F., Aller, H. D., et al. 2015, *ApJL*, **810**, L9
- Lister, M. L., & Homan, D. C. 2005, *AJ*, **130**, 1389
- Lister, M. L., & Marscher, A. P. 1997, *ApJ*, **476**, 572
- Lister, M. L., Aller, H. D., Aller, M. F., et al. 2009a, *AJ*, **137**, 3718
- Lister, M. L., Cohen, M. H., Homan, D. C., et al. 2009b, *AJ*, **138**, 1874
- Lister, M. L., Aller, M., Aller, H., et al. 2011, *ApJ*, **742**, 27
- Lister, M. L., Aller, M. F., Aller, H. D., et al. 2013, *AJ*, **146**, 120
- Ludke, E., Garrington, S. T., Spencer, R. E., et al. 1998, *MNRAS*, **299**, 467
- Lynds, C. R. 1967, *ApJ*, **147**, 837
- Marcha, M. J. M., Browne, I. W. A., Impey, C. D., & Smith, P. S. 1996, *MNRAS*, **281**, 425
- Marr, J. M., Taylor, G. B., & Crawford, F., III 2001, *ApJ*, **550**, 160
- Marziani, P., Sulentic, J. W., Dultzin-Hacyan, D., Calvani, M., & Moles, M. 1996, *ApJS*, **104**, 37
- Marziani, P., Sulentic, J. W., Zamanov, R., et al. 2003, *ApJS*, **145**, 199
- McIntosh, D. H., Rieke, M. J., Rix, H.-W., Foltz, C. B., & Weymann, R. J. 1999, *ApJ*, **514**, 40
- Meisner, A. M., & Romani, R. W. 2010, *ApJ*, **712**, 14
- Michel, A., & Huchra, J. 1988, *PASP*, **100**, 1423
- Nilsson, K., Pursimo, T., Villforth, C., et al. 2012, *A&A*, **547**, A1
- Osmer, P. S., Porter, A. C., & Green, R. F. 1994, *ApJ*, **436**, 678
- Owen, F. N., Ledlow, M. J., & Keel, W. C. 1995, *AJ*, **109**, 14
- Owen, F. N., Ledlow, M. J., Morrison, G. E., & Hill, J. M. 1997, *ApJL*, **488**, L15
- Paiano, S., Landoni, M., Falomo, R., Scarpa, R., & Treves, A. 2016, *MNRAS*, **458**, 2836
- Perlman, E. S., Padovani, P., Giommi, P., et al. 1998, *AJ*, **115**, 1253
- Perlman, E. S., Stocke, J. T., Carilli, C. L., et al. 2002, *AJ*, **124**, 2401
- Perlman, E. S., Stocke, J. T., Schachter, J. F., et al. 1996, *ApJS*, **104**, 251
- Peterson, B. A., Wright, A. E., Jauncey, D. L., & Condon, J. J. 1979, *ApJ*, **232**, 400
- Pietsch, W., Bischoff, K., Boller, T., et al. 1998, *A&A*, **333**, 48
- Piner, B. G., Pant, N., & Edwards, P. G. 2008, *ApJ*, **678**, 64
- Pita, S., Goldoni, P., Boisson, C., et al. 2014, *A&A*, **565**, A12
- Pravdo, S. H., & Marshall, F. E. 1984, *ApJ*, **281**, 570
- Punsly, B. 1999, *ApJ*, **516**, 141
- Pushkarev, A. B., Hovatta, T., Kovalev, Y. Y., et al. 2012, *A&A*, **545**, A113
- Pushkarev, A. B., & Kovalev, Y. Y. 2015, *MNRAS*, **452**, 4274
- Pushkarev, A. B., Kovalev, Y. Y., Lister, M. L., et al. 2013, *A&A*, **555**, A80
- Rau, A., Schady, P., Greiner, J., et al. 2012, *A&A*, **538**, A26
- Rector, T. A., & Stocke, J. T. 2001, *AJ*, **122**, 565
- Richards, G. T., Myers, A. D., Gray, A. G., et al. 2009, *ApJS*, **180**, 67
- Rovero, A. C., Muriel, H., Donzelli, C., & Pichel, A. 2016, *A&A*, **589**, A92
- Sargent, W. L. W. 1970, *ApJ*, **160**, 405
- Savage, A., Browne, I. W. A., & Bolton, J. G. 1976, *MNRAS*, **177**, 77
- Sbarufatti, B., Treves, A., & Falomo, R. 2005a, *ApJ*, **635**, 173
- Sbarufatti, B., Treves, A., Falomo, R., et al. 2005b, *AJ*, **129**, 559
- Sbarufatti, B., Treves, A., Falomo, R., et al. 2006, *AJ*, **132**, 1
- Schachter, J. F., Stocke, J. T., Perlman, E., et al. 1993, *ApJ*, **412**, 541
- Schmidt, M. 1977, *ApJ*, **217**, 358
- Schneider, D. P., Schmidt, M., & Gunn, J. E. 1999, *AJ*, **117**, 40
- Schneider, D. P., Richards, G. T., Fan, X., et al. 2002, *AJ*, **123**, 567
- Schneider, D. P., Fan, X., Hall, P. B., et al. 2003, *AJ*, **126**, 2579
- Schneider, D. P., Hall, P. B., Richards, G. T., et al. 2005, *AJ*, **130**, 367
- Schneider, D. P., Hall, P. B., Richards, G. T., et al. 2007, *AJ*, **134**, 102
- Schneider, D. P., Richards, G. T., Hall, P. B., et al. 2010a, *AJ*, **139**, 2360

- Schneider, D. P., Richards, G. T., Hall, P. B., et al. 2010b, *AJ*, **139**, 2360
- Schramm, K.-J., Borgeest, U., Kuehl, D., et al. 1994, *A&AS*, **106**, 349
- Searle, L., & Bolton, J. G. 1968, *ApJL*, **154**, L101
- Shaw, M. S., Filippenko, A. V., Romani, R. W., Cenko, S. B., & Li, W. 2013a, *AJ*, **146**, 127
- Shaw, M. S., Romani, R. W., Healey, S. E., et al. 2009, *ApJ*, **704**, 477
- Shaw, M. S., Romani, R. W., Cotter, G., et al. 2012, *ApJ*, **748**, 49
- Shaw, M. S., Romani, R. W., Cotter, G., et al. 2013b, *ApJ*, **764**, 135
- Shepherd, M. C. 1997, in *ASP Conf. Ser.* 125, *Astronomical Data Analysis Software and Systems VI*, ed. G. Hunt, & H. E. Payne (San Francisco, CA: ASP), 77
- Small, T. A., Sargent, W. L. W., & Steidel, C. C. 1997, *AJ*, **114**, 2254
- Smith, H. E., Burbidge, E. M., Baldwin, J. A., et al. 1977, *ApJ*, **215**, 427
- Smith, R. J., Lucey, J. R., Hudson, M. J., Schlegel, D. J., & Davies, R. L. 2000, *MNRAS*, **313**, 469
- Sowards-Emmerd, D., Romani, R. W., & Michelson, P. F. 2003, *ApJ*, **590**, 109
- Sowards-Emmerd, D., Romani, R. W., Michelson, P. F., Healey, S. E., & Nolan, P. L. 2005, *ApJ*, **626**, 95
- Spinrad, H., Marr, J., Aguilar, L., & Djorgovski, S. 1985, *PASP*, **97**, 932
- Stadnik, M., & Romani, R. W. 2014, *ApJ*, **784**, 151
- Steidel, C. C., Adelberger, K. L., Shapley, A. E., et al. 2003, *ApJ*, **592**, 728
- Steidel, C. C., & Sargent, W. L. W. 1991, *ApJ*, **382**, 433
- Stickel, M., Fried, J. W., & Kuehr, H. 1988, *A&A*, **191**, L16
- Stickel, M., Fried, J. W., & Kuehr, H. 1989, *A&AS*, **80**, 103
- Stickel, M., Fried, J. W., & Kuehr, H. 1993a, *A&AS*, **98**, 393
- Stickel, M., & Kuehr, H. 1993, *A&AS*, **100**, 395
- Stickel, M., & Kuehr, H. 1996, *A&AS*, **115**, 11
- Stickel, M., Kuehr, H., & Fried, J. W. 1993b, *A&AS*, **97**, 483
- Stickel, M., & Kuhr, H. 1993, *A&AS*, **101**, 521
- Stickel, M., Meisenheimer, K., & Kuehr, H. 1994, *A&AS*, **105**, 211
- Stoeckle, J. T., Wurtz, R., Wang, Q., Elston, R., & Jannuzi, B. T. 1992, *ApJL*, **400**, L17
- Strauss, M. A., Huchra, J. P., Davis, M., et al. 1992, *ApJS*, **83**, 29
- Strittmatter, P. A., Carswell, R. F., Gilbert, G., & Burbidge, E. M. 1974, *ApJ*, **190**, 509
- Tadhunter, C. N., Morganti, R., di Serego-Alighieri, S., Fosbury, R. A. E., & Danziger, I. J. 1993, *MNRAS*, **263**, 999
- Thompson, D. J., Djorgovski, S., & de Carvalho, R. 1990, *PASP*, **102**, 1235
- Thompson, D. J., Djorgovski, S., Vigotti, M., & Gruelf, G. 1992, *ApJS*, **81**, 1
- Titov, O., Jauncey, D. L., Johnston, H. M., Hunstead, R. W., & Christensen, L. 2011, *AJ*, **142**, 165
- Ulrich, M.-H., Kinman, T. D., Lynds, C. R., Rieke, G. H., & Ekers, R. D. 1975, *ApJ*, **198**, 261
- Urry, C. M., & Padovani, P. 1995, *PASP*, **107**, 803
- Vandenbroucke, J., Buehler, R., Ajello, M., et al. 2010, *ApJL*, **718**, L166
- Vermeulen, R. C., & Cohen, M. H. 1994, *ApJ*, **430**, 467
- Vermeulen, R. C., Ogle, P. M., Tran, H. D., et al. 1995, *ApJL*, **452**, L5
- Vermeulen, R. C., Taylor, G. B., Readhead, A. C. S., & Browne, I. W. A. 1996, *AJ*, **111**, 1013
- Véron-Cetty, M. P., & Véron, P. 2000, *A&ARv*, **10**, 81
- Walsh, D., Beckers, J. M., Carswell, R. F., & Weymann, R. J. 1984, *MNRAS*, **211**, 105
- Walsh, D., & Carswell, R. F. 1982, *MNRAS*, **200**, 191
- Wegner, G., Bernardi, M., Willmer, C. N. A., et al. 2003, *AJ*, **126**, 2268
- White, G. L., Jauncey, D. L., Wright, A. E., et al. 1988, *ApJ*, **327**, 561
- White, R. L., Becker, R. H., Gregg, M. D., et al. 2000, *ApJS*, **126**, 133
- Wilkes, B. J. 1986, *MNRAS*, **218**, 331
- Wilkes, B. J., Wright, A. E., Jauncey, D. L., & Peterson, B. A. 1983, *PASAu*, **5**, 2
- Wills, D., & Lynds, R. 1978, *ApJS*, **36**, 317
- Wills, D., & Wills, B. J. 1974, *ApJ*, **190**, 271
- Wills, D., & Wills, B. J. 1976, *ApJS*, **31**, 143
- Wright, A. E., Ables, J. G., & Allen, D. A. 1983, *MNRAS*, **205**, 793
- Wright, A. E., Peterson, B. A., Jauncey, D. L., & Condon, J. J. 1979, *ApJ*, **229**, 73
- Xu, W., Lawrence, C. R., Readhead, A. C. S., & Pearson, T. J. 1994, *AJ*, **108**, 395
- Yao, S., Yuan, W., Zhou, H., et al. 2015, *MNRAS*, **454**, L16
- York, D. G., Adelman, J., Anderson, J. E., Jr., et al. 2000, *AJ*, **120**, 1579
- Young, P., Sargent, W. L. W., & Boksenberg, A. 1982, *ApJS*, **48**, 455
- Zensus, J. A., Ros, E., Kellermann, K. I., et al. 2002, *AJ*, **124**, 662